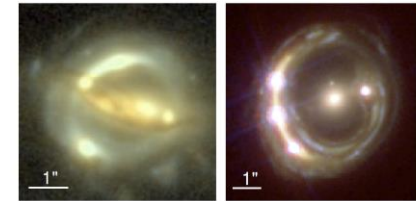
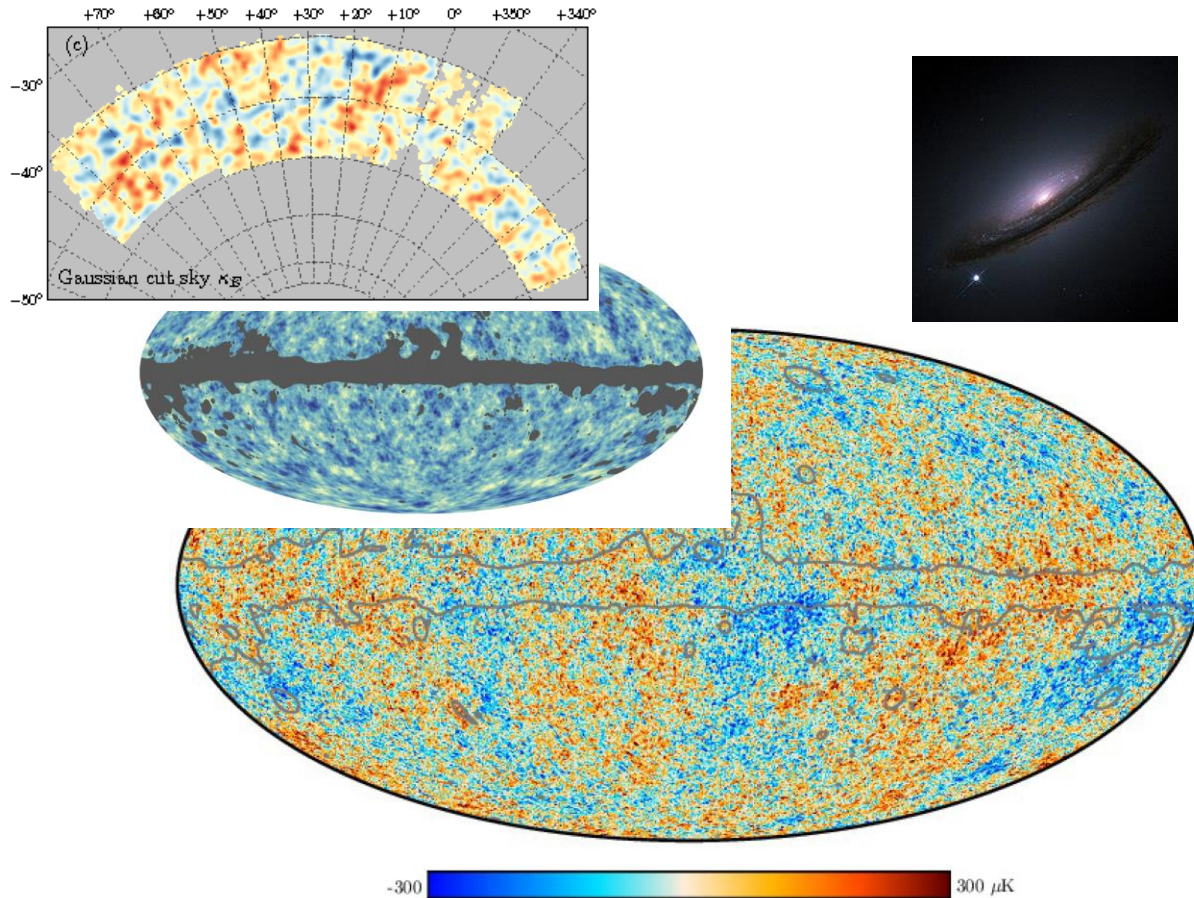
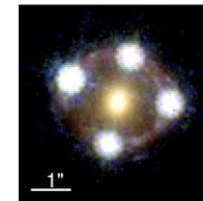


Cosmic concordance, tensions and lensing

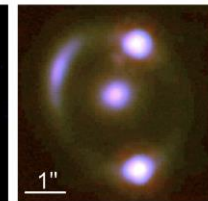


(a) B1608+656

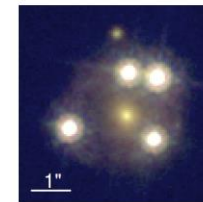
(b) RXJ1131-1231



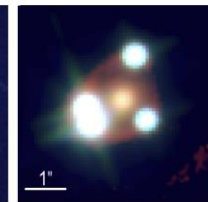
(c) HE 0435-1223



(d) SDSS 1206+4332

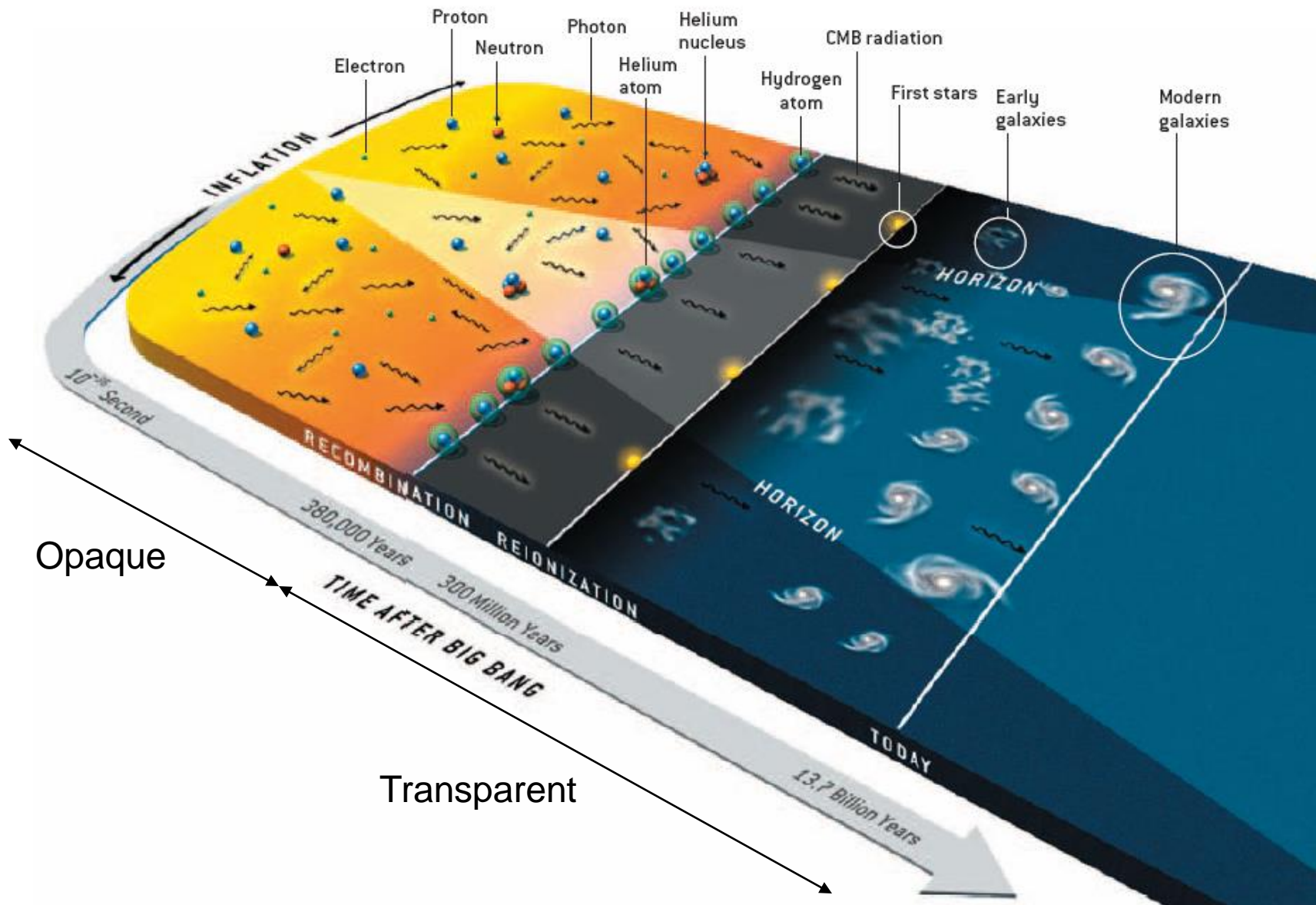


(e) WFI2033-4723

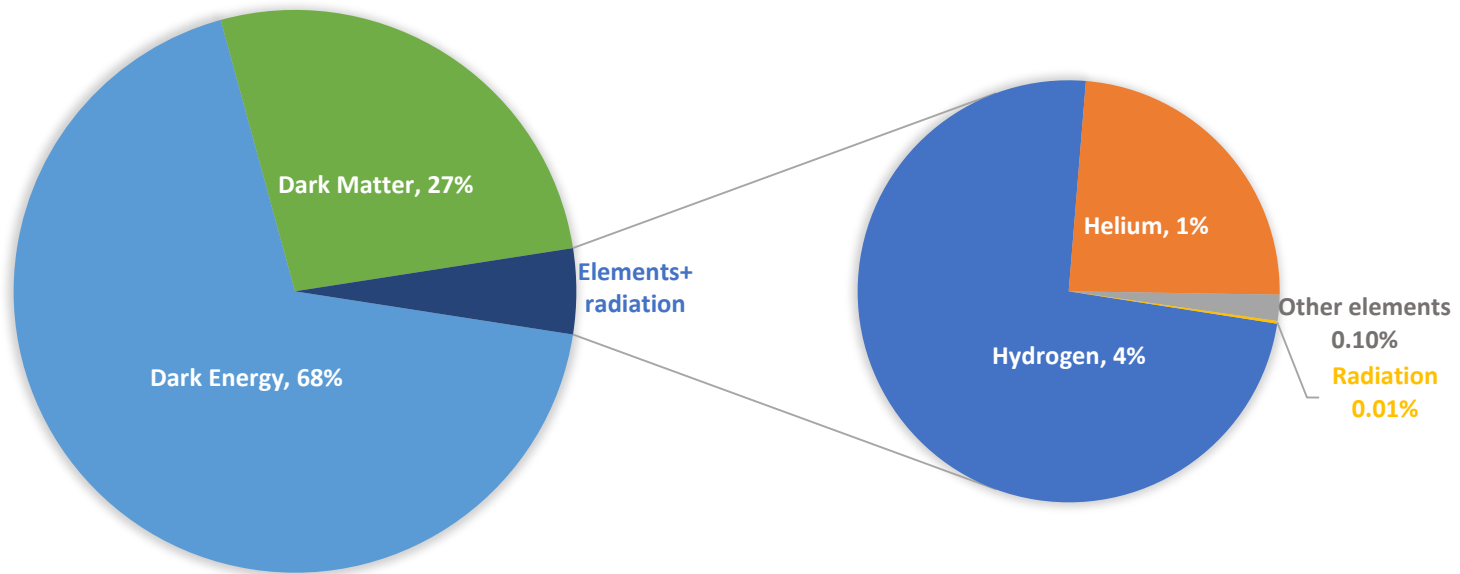


(f) PG 1115+080

Evolution in the standard cosmology



Contents of the Universe today

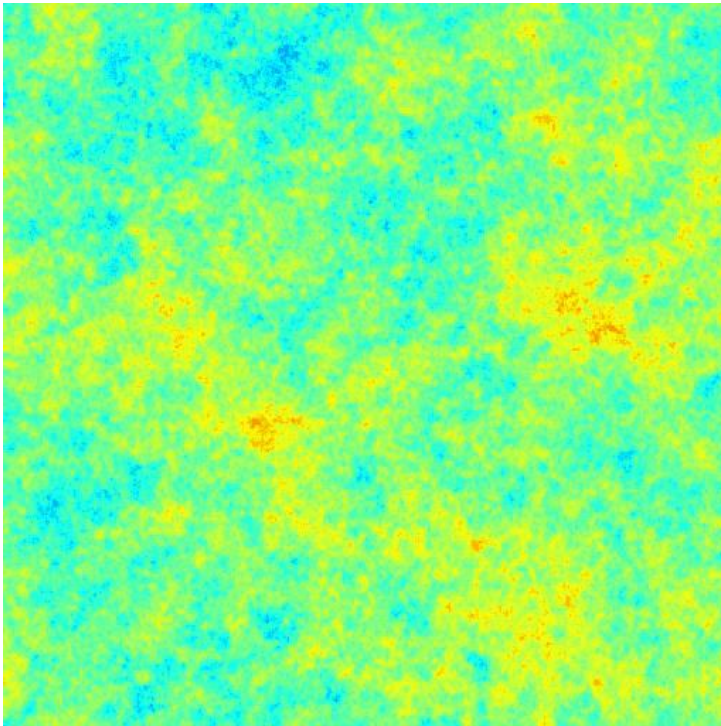


Define and test perturbatively- Λ CDM model:

- Photons (CMB temperature today ~ 2.7255 K)
 - 3 active neutrinos, assuming minimal mass hierarchy with $\sum m_\nu = 0.06$ eV
 - Standard model baryons (taken to include electrons etc), density $\Omega_b h^2$
 - Cold (pressureless) non-interacting and stable matter (CDM), density $\Omega_c h^2$
 - Cosmological constant, giving a flat universe with $\Omega_K = 0$
 - Reionization parameterized by a single effective optical depth τ
 - Gaussian adiabatic primordial curvature perturbations with power spectrum $P_R = A_s \left(\frac{k}{k_*} \right)^{n_s-1}$
- Remaining free parameter is $H_0 = 100h \text{ km s}^{-1} \text{ Mpc}^{-1}$ (or $\Omega_\Lambda, \Omega_m, \theta_*, \dots$)*

Perturbation evolution

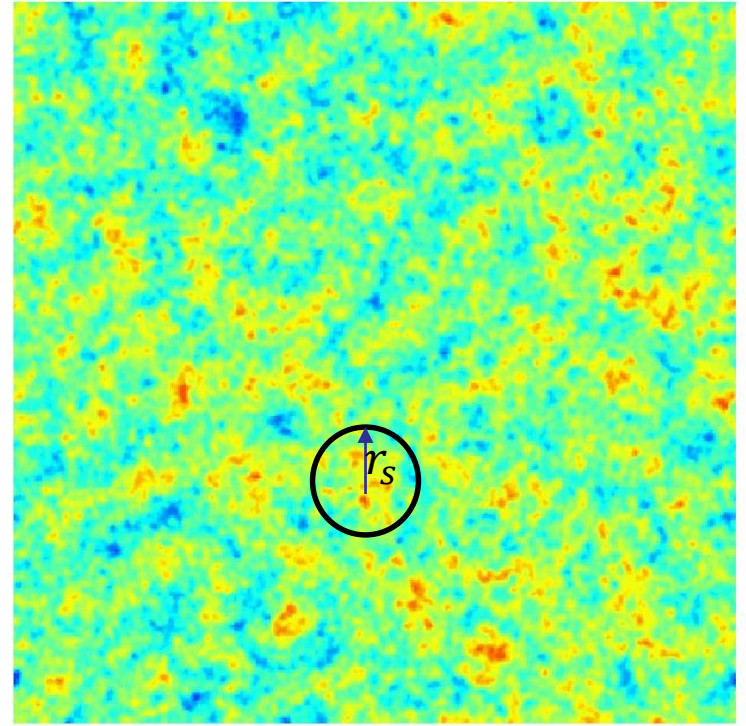
Perturbations: start of hot big bang



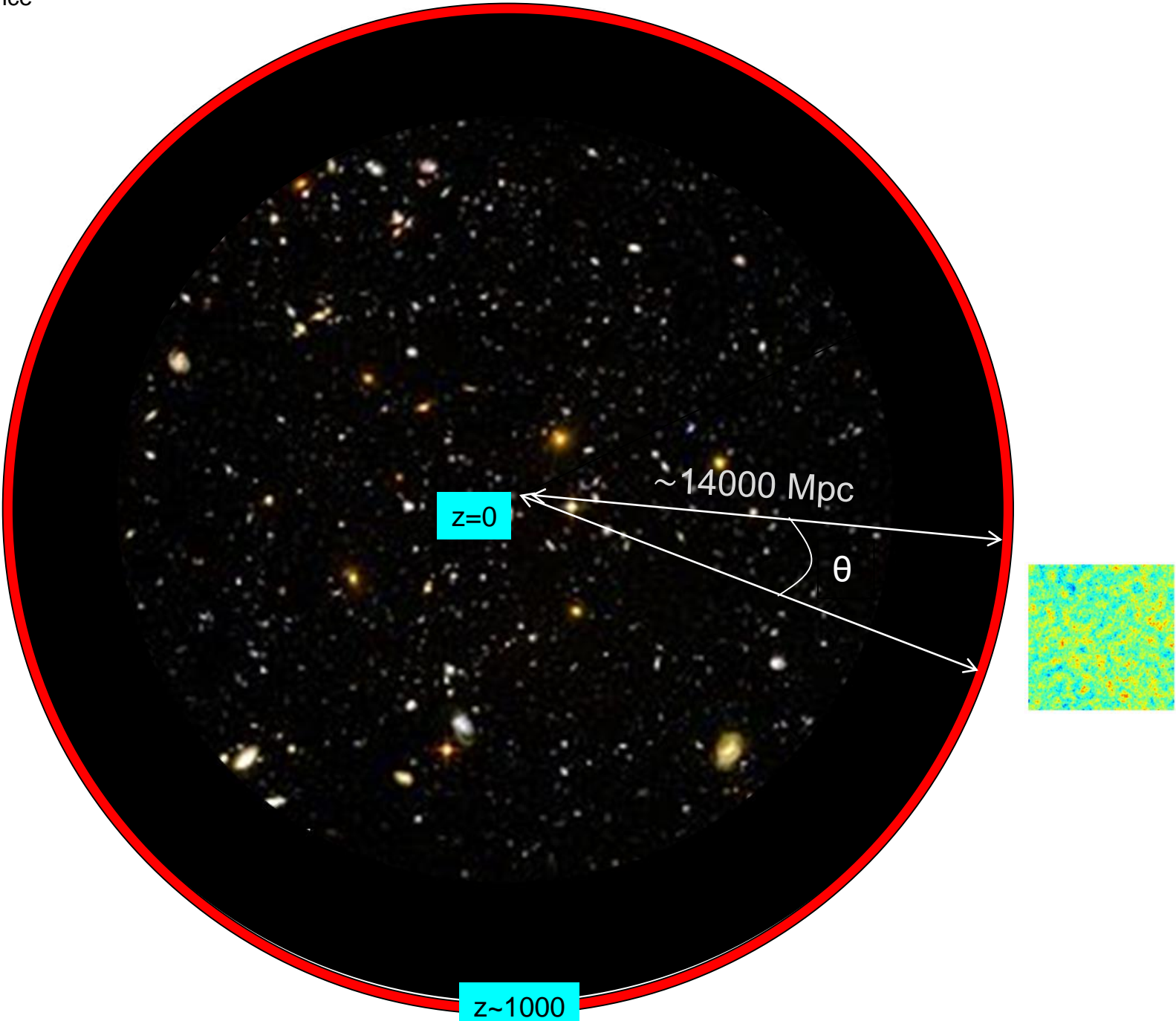
gravity+
pressure+
diffusion



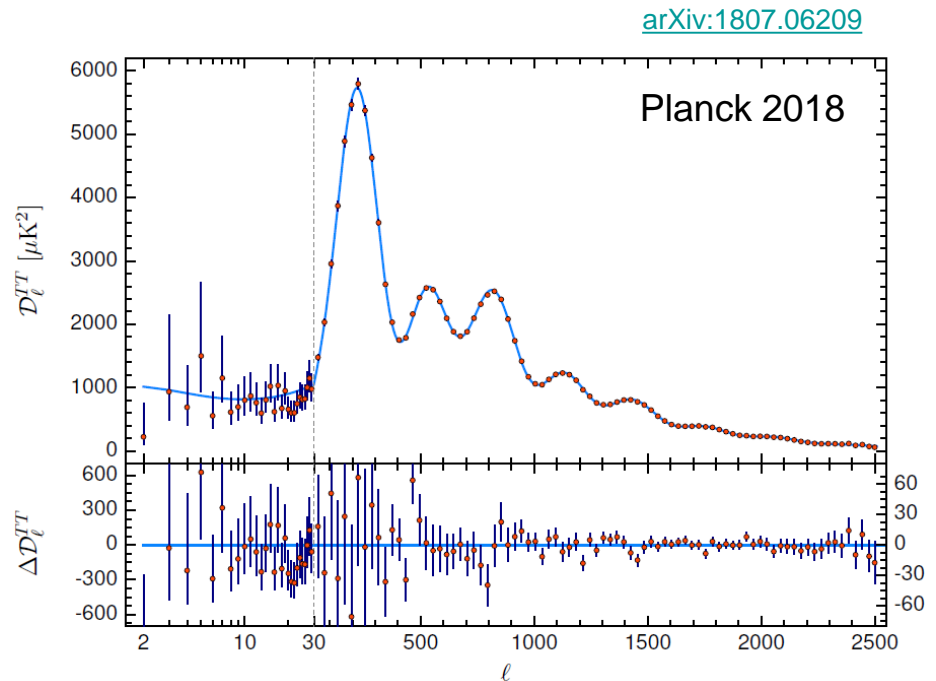
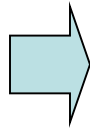
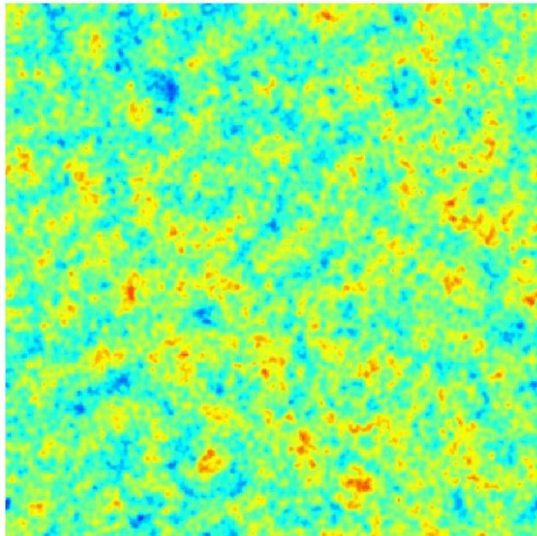
Perturbations: Last scattering surface



In comoving distance



Observed CMB power spectrum



Observations
(10^{-5} perturbations)



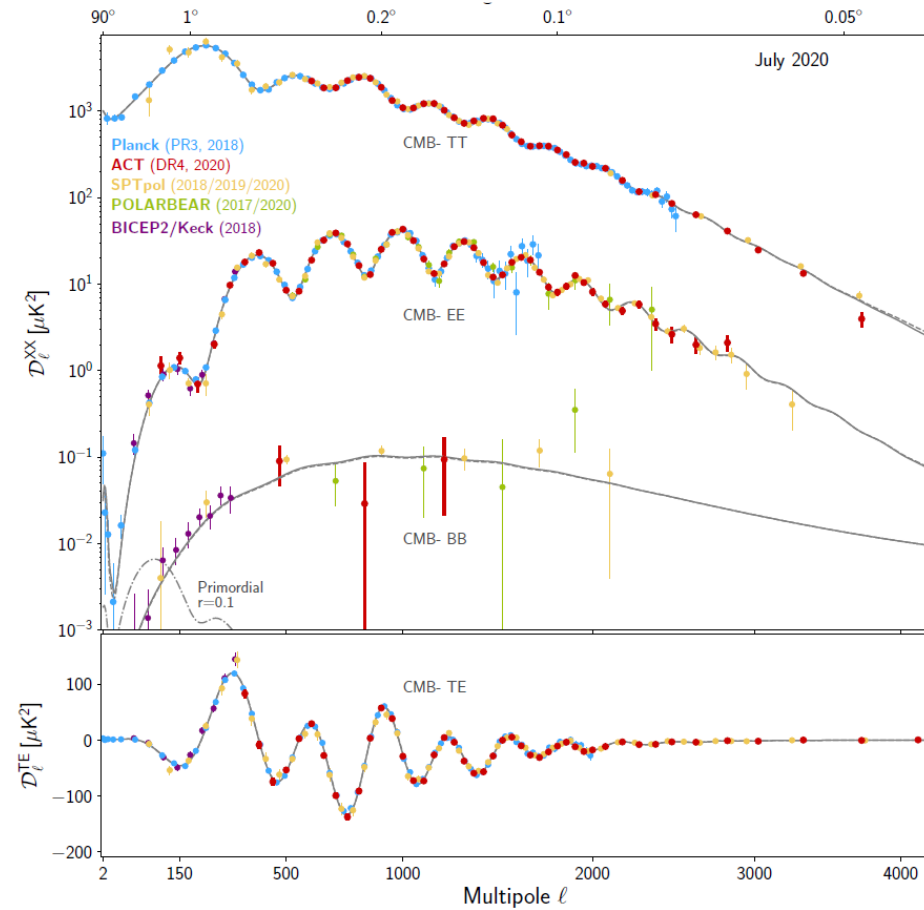
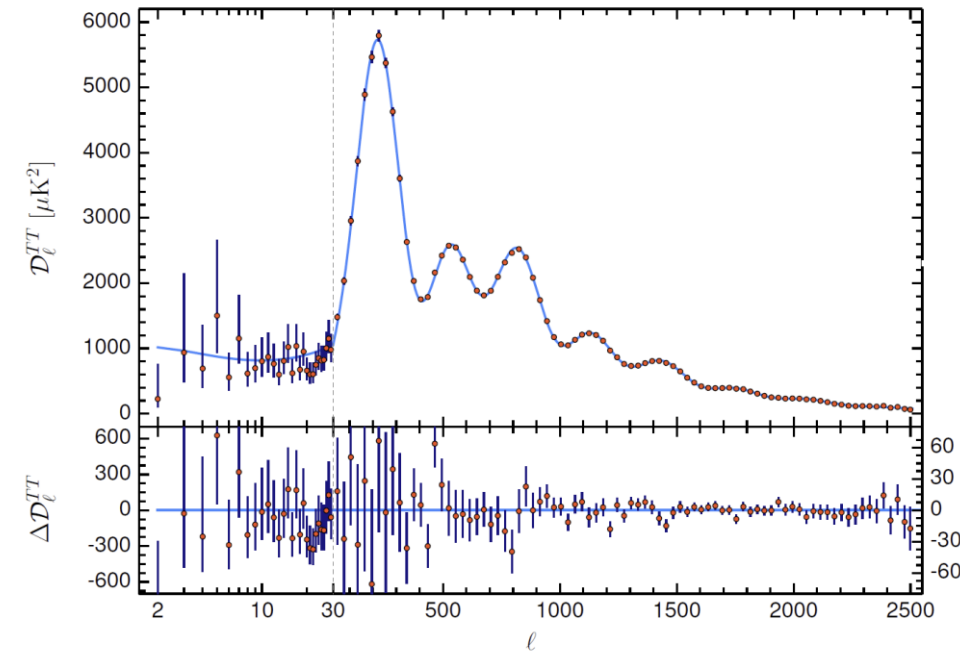
Assume model, constrain parameters
- test constancy with other probes

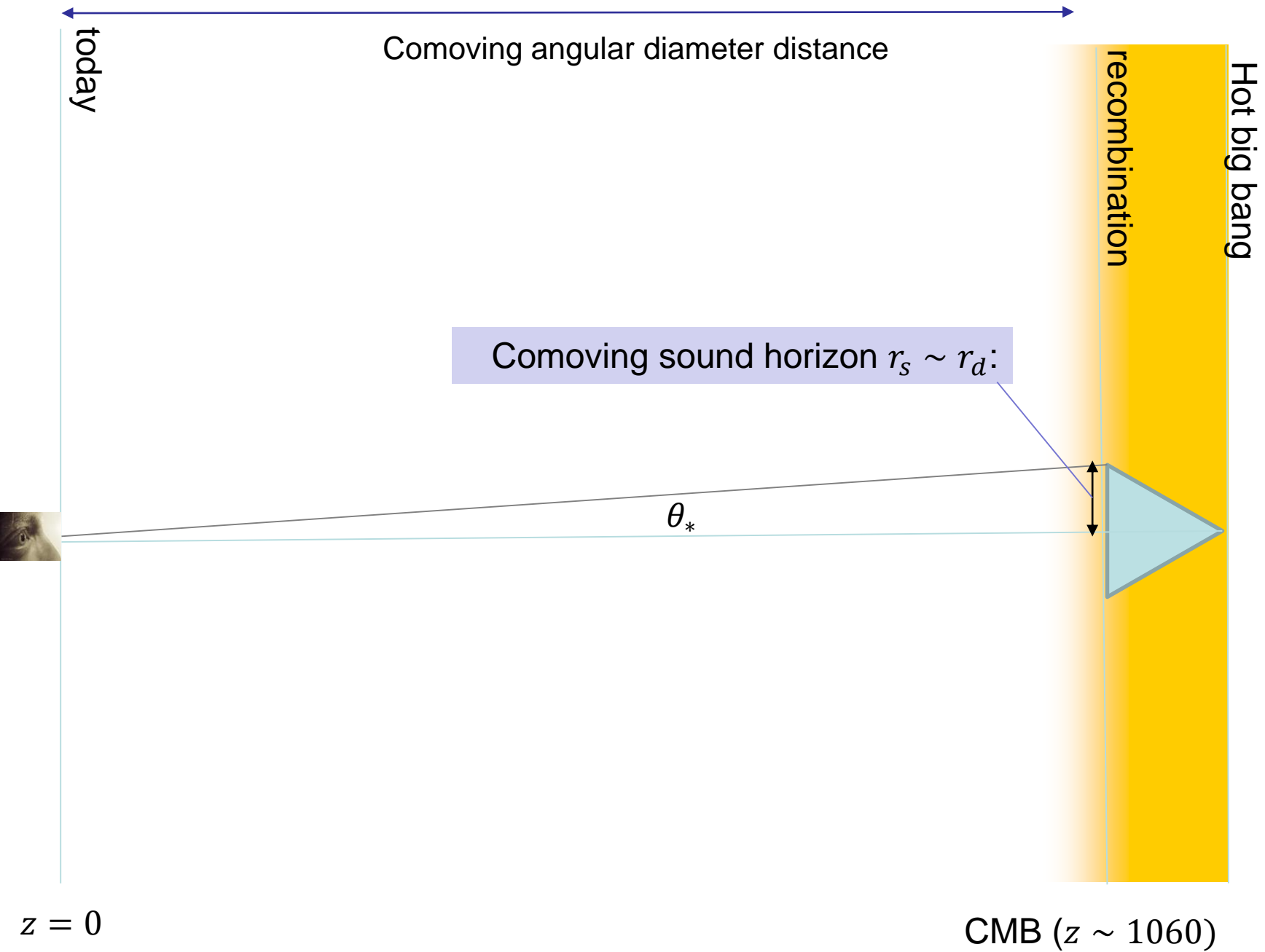
Linear perturbation theory very accurate: given a model, can calculate to high precision

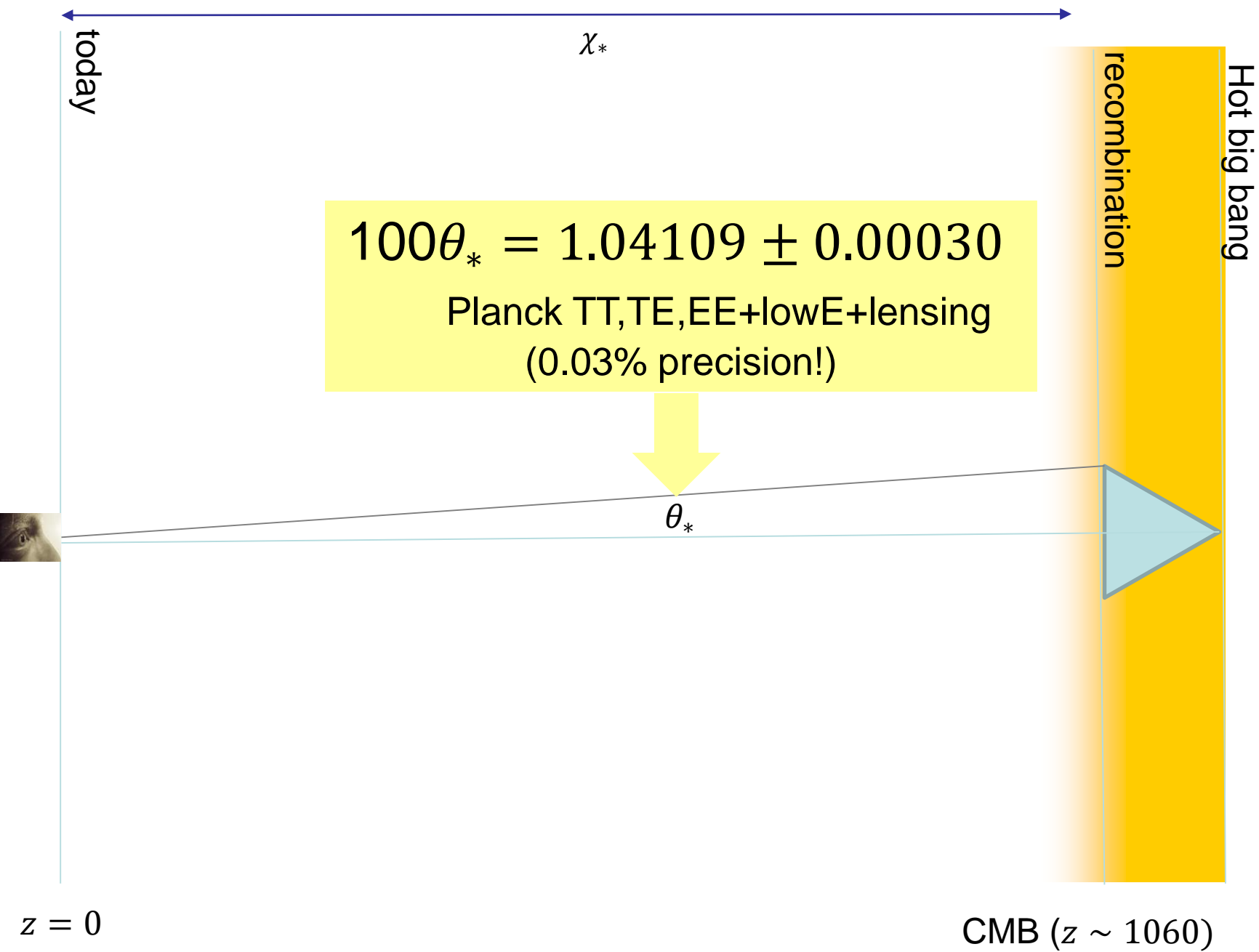
Cosmic Microwave Background power spectrum fits to Λ CDM

Choi et al. (2007.07289)

Planck 2018 (arXiv: 1806.06209)

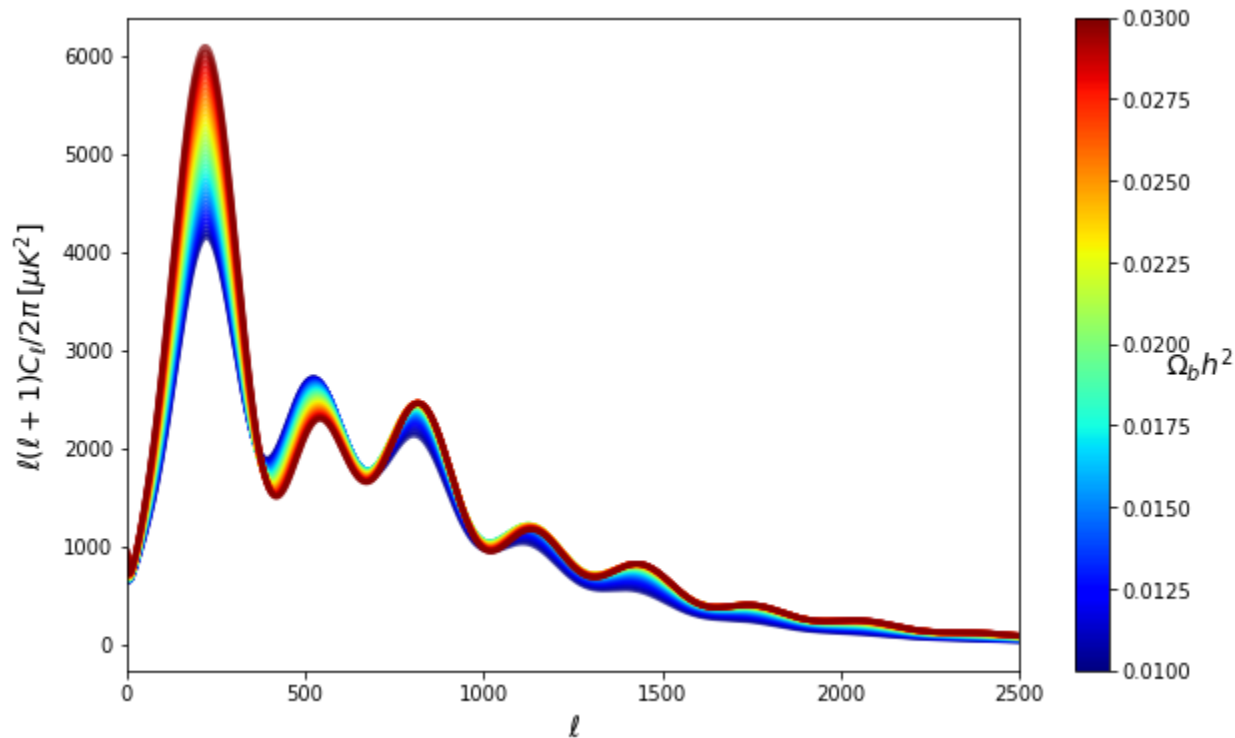






Λ CDM baryon density at fixed θ_* , $\Omega_m h^2$

(baryons deepen overdensity compressions: enhance odd peaks of spectrum)



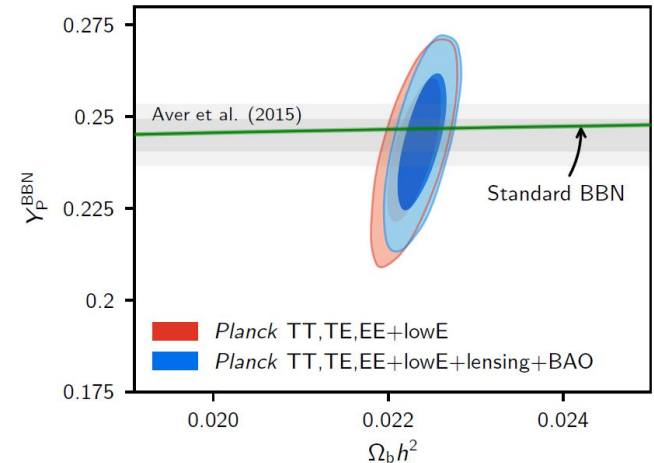
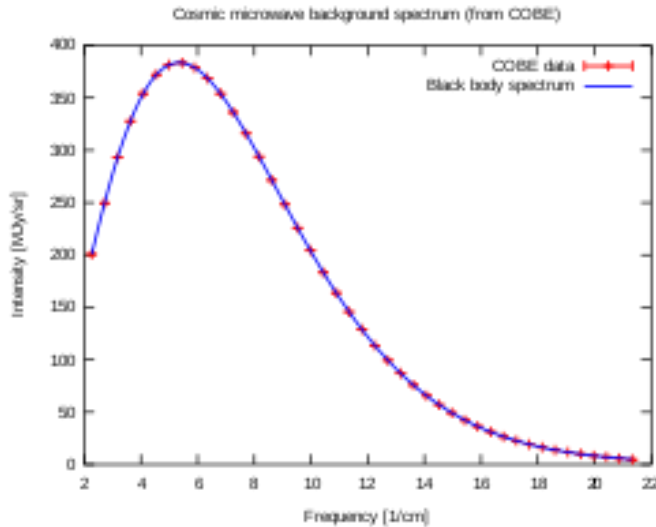
Odd/even height ratio distinctive and quite robust:

$$\Omega_b h^2 = 0.0224 \pm 0.0002$$

Consistency with standard Big-Bang Nucleosynthesis

arXiv: 1806.06209

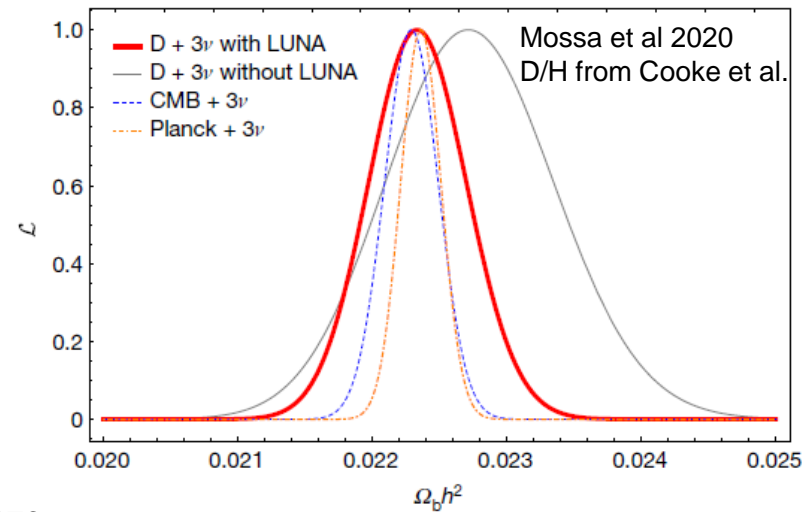
COBE measured $T_{\text{CMB}} \sim 2.7255 \text{ K}$



BUT: Lithium problem remains around 5σ

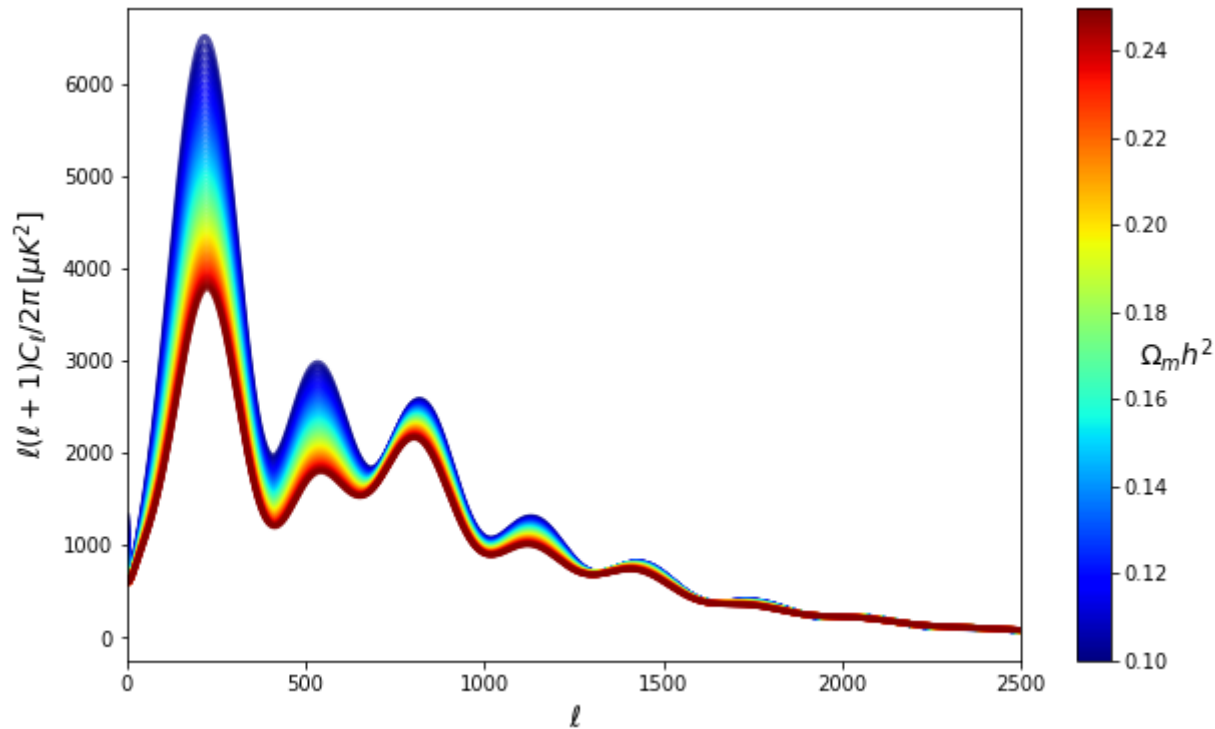
Measured: ${}^7\text{Li}/\text{H} = (1.58 \pm 0.35) \times 10^{-10}$ arXiv: 1505.01076

Prediction: ${}^7\text{Li}/\text{H} = 4.5 \times 10^{-10}$



Λ CDM matter density at fixed θ_* , $\Omega_b h^2$

(more matter *lowers* amplitude for modes that enter horizon in matter domination)



Can be partly compensated by changing initial power A_s, n_s and foregrounds.
But detailed shape is still quite distinctive and robust:

$$\Omega_m h^2 = 0.143 \pm 0.001$$

← today →

Assume baryons, CDM, photons, 3 neutrinos
Know T_{CMB} , peaks measure $\Omega_m h^2, \Omega_b h^2$

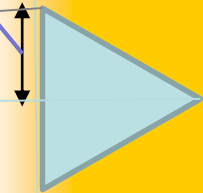
⇒ comoving sound horizon:

$$r_s \approx \int_0^{t_*} \frac{c_s dt}{a} \sim (144.4 \pm 0.3) \text{ Mpc}$$

recombination

Hot big bang

θ_*



$z = 0$

CMB ($z \sim 1060$)

$r_s, \theta_* \Rightarrow$ Comoving radial distance $\chi_* \sim (13.87 \pm 0.03)$ Gpc

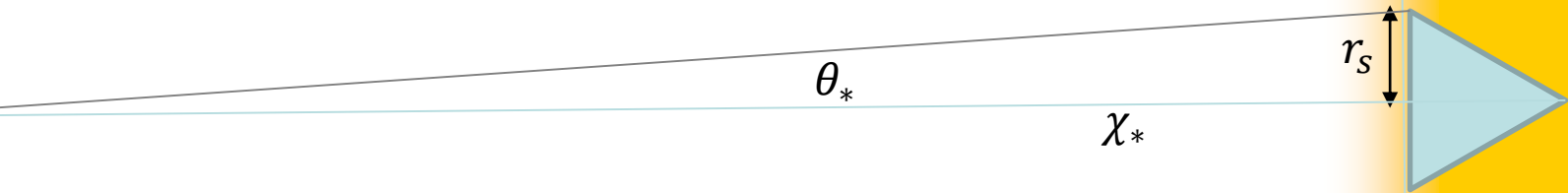
$$\begin{aligned}\chi_* &= \int \left(\frac{cdt}{a} \right) \\ &= \int \left(\frac{da}{a^2 H} \right) \approx \int \frac{da}{\sqrt{a\Omega_m H_0^2 + a^4\Omega_\Lambda H_0^2}}\end{aligned}$$

$$\Omega_\Lambda H_0^2 = H_0^2 - \Omega_m H_0^2 \text{ and know } \Omega_m h^2 \Rightarrow H_0$$

today

recombination

Hot big bang



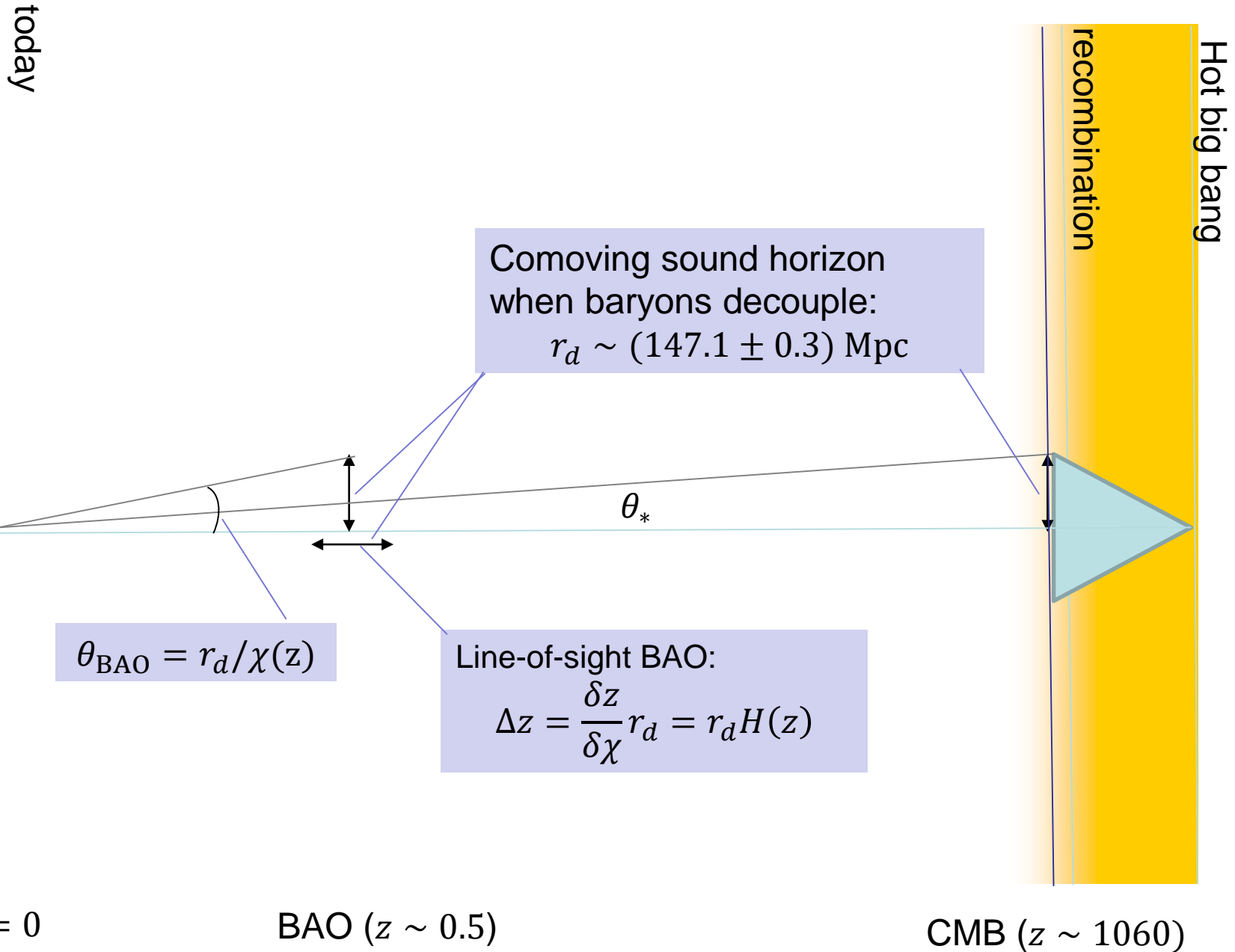
$$\Rightarrow H_0 = (67.3 \pm 0.6) \text{ km s}^{-1} \text{ Mpc}^{-1}$$

(Planck, confirmed by ACT)

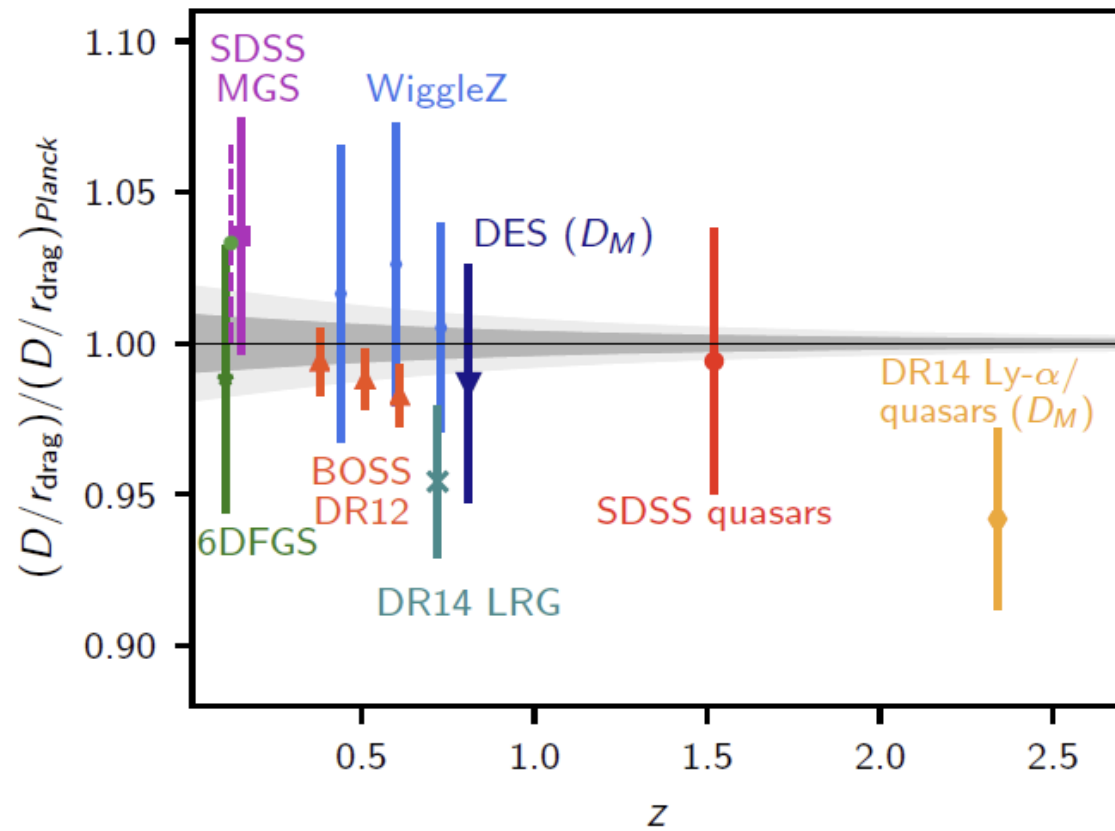
$z = 0$

CMB ($z \sim 1060$)

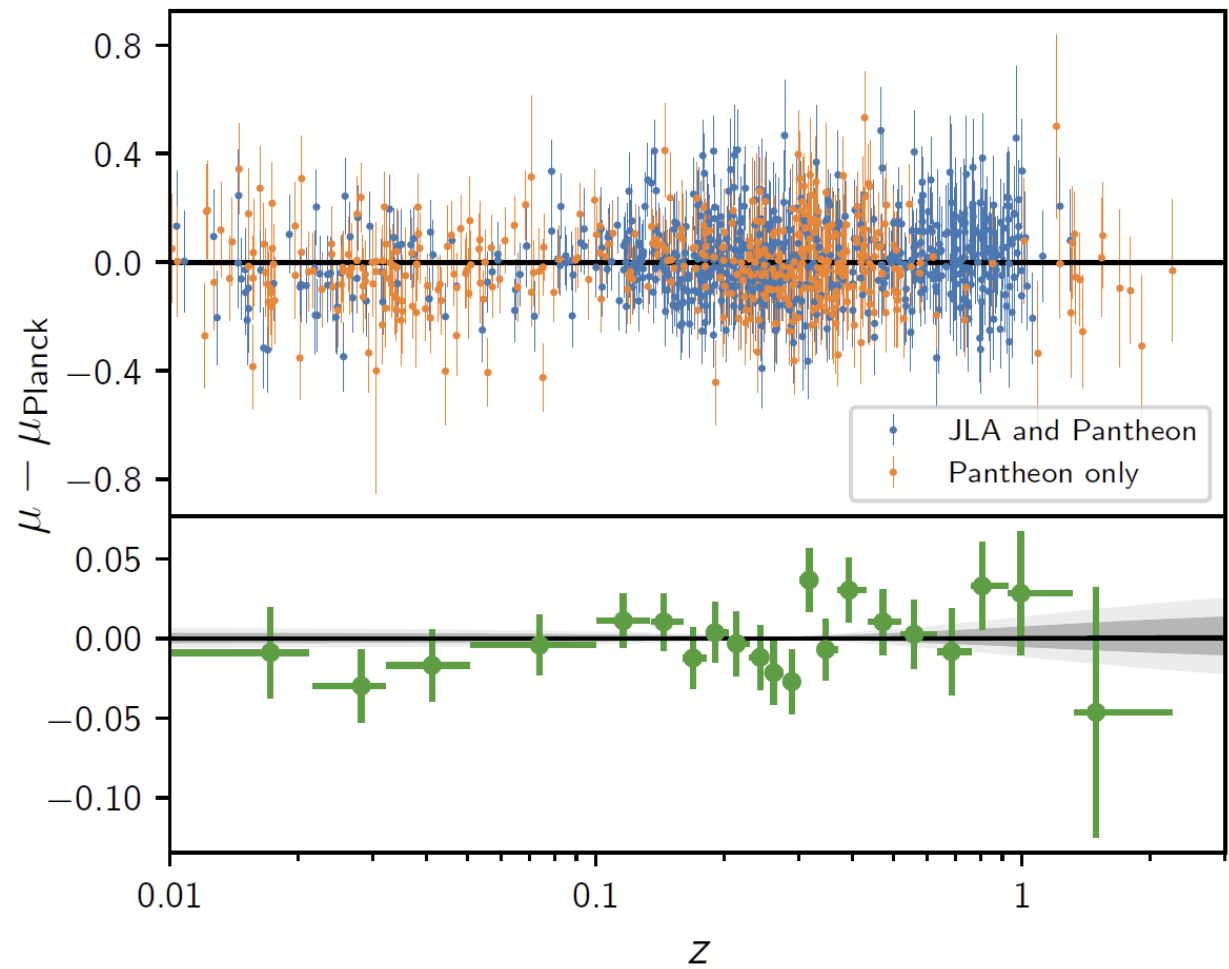
CMB and BAO consistency in Λ CDM



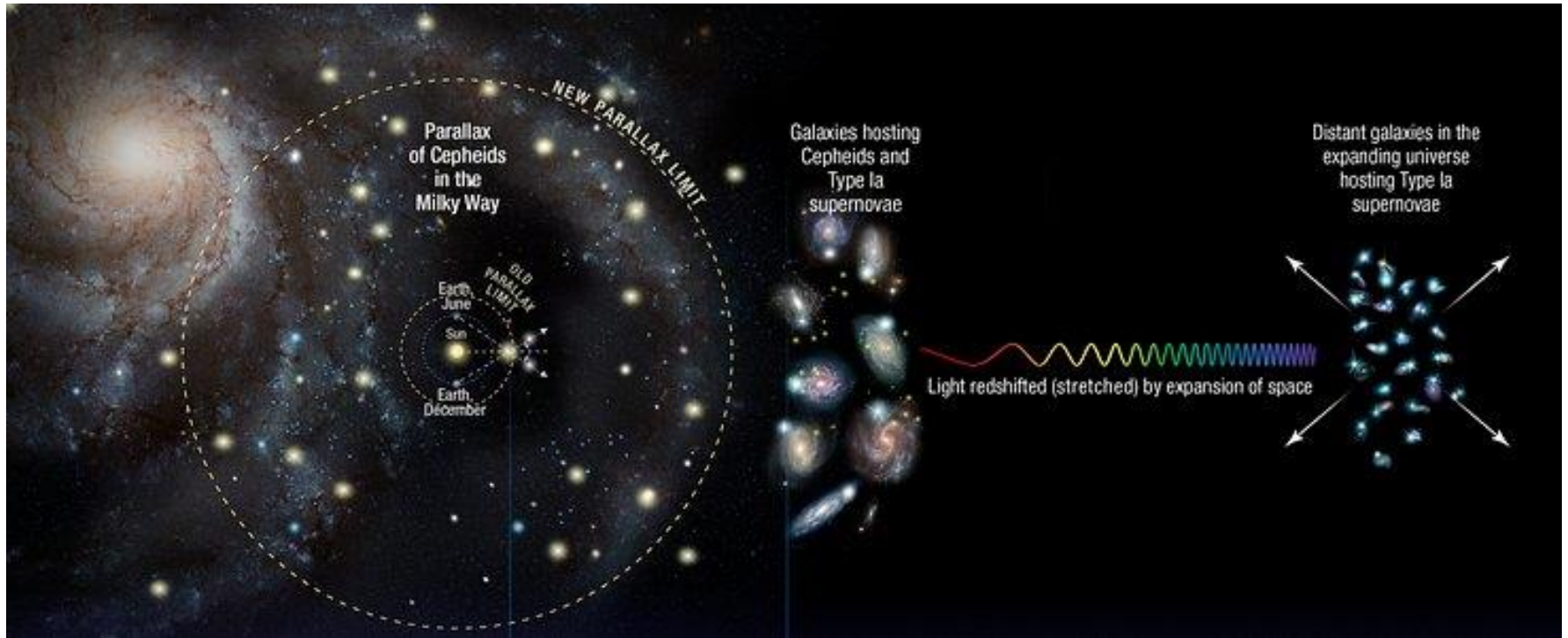
Baryon Acoustic Oscillation (BAO) concordance



Supernovae constrain redshift evolution (as standardizable candles, measure d_L)



H_0 from local distance ladder



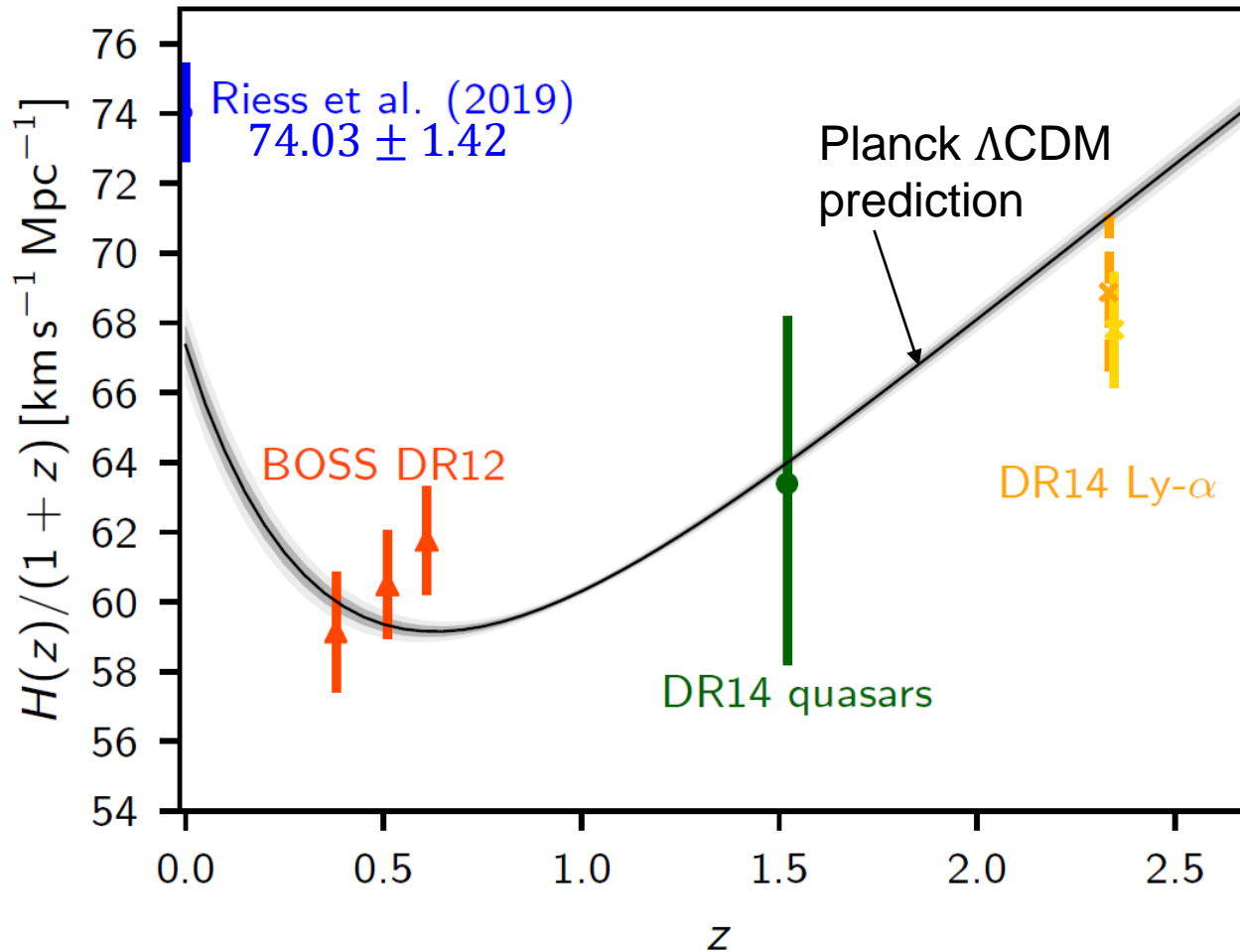
<https://www.spacetelescope.org/news/heic1611/>

Parallax+cepheids+SN

$$H_0 = 74.03 \pm 1.42 \text{ km s}^{-1} \text{Mpc}^{-1}$$

Riess et al. arXiv: 1903.07603

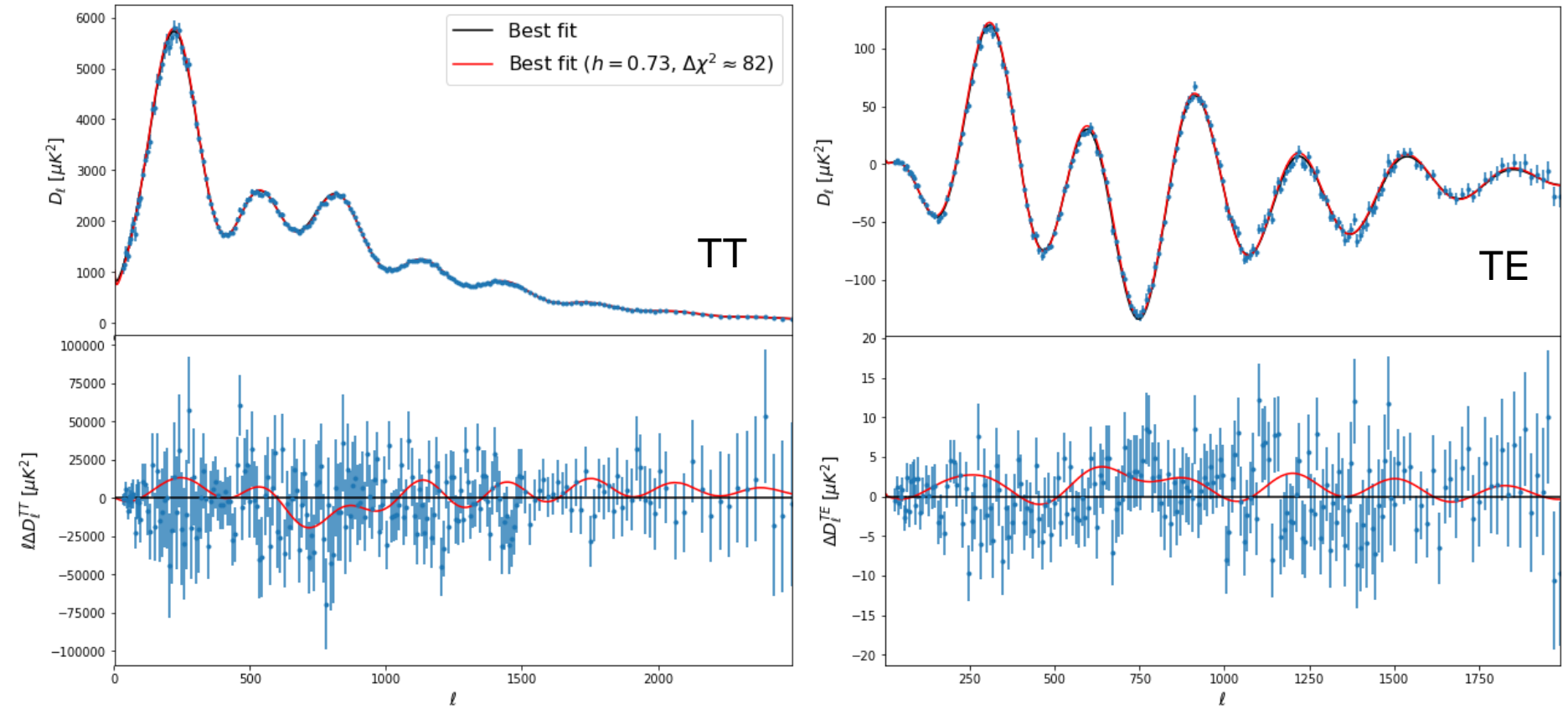
The Hubble discrepancy assuming Λ CDM and Planck sound horizon r_d



A $4.4\sigma \sim 10\%$ discrepancy between local and CMB-inferred Λ CDM H_0 ?

Model fits

LCDM best-fits: $H_0 = 67.3$ ($n_s = 0.966$, $\Omega_m = 0.32$, $\Omega_m h^2 = 0.143$)
vs. best fit for $H_0 = 73.0$ ($n_s = 0.995$, $\Omega_m = 0.25$, $\Omega_m h^2 = 0.132$)



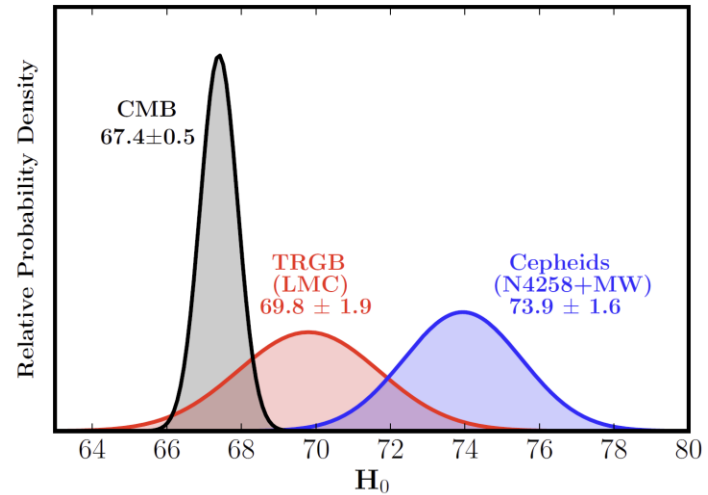
Other local Hubble parameter measurements

Forward distance ladder

Tip of the red giant branch

$$H_0 = 69.8 \pm 1.9 \text{ km s}^{-1} \text{Mpc}^{-1}$$

Freedman et al. arXiv:1907.05922, 2002.01550



+ other several results using other local calibrators, all giving high H_0

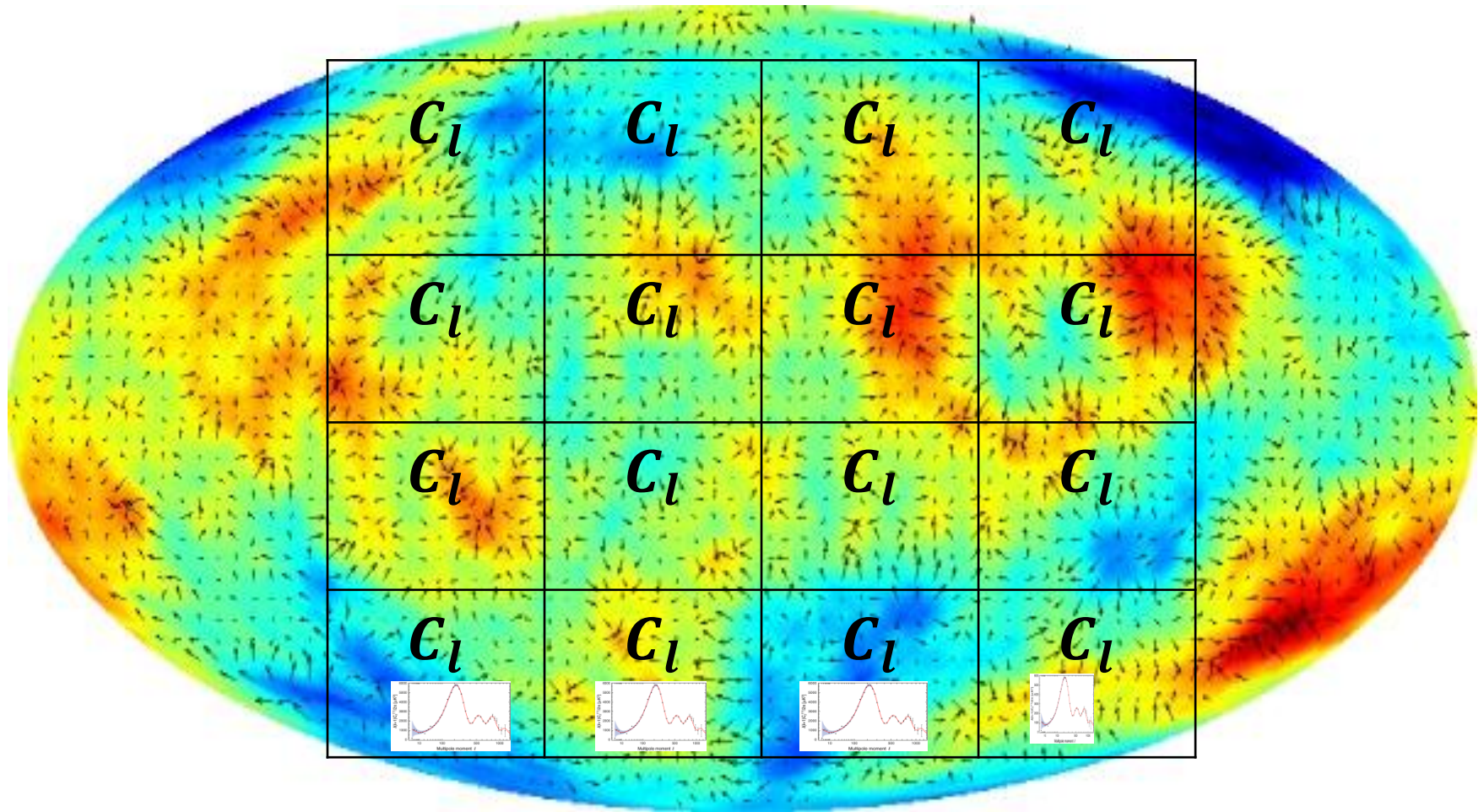
Planck Lensing Reconstruction

Planck Lensing 2018 arXiv:1807.06210

Lead by Julien Carron at Sussex



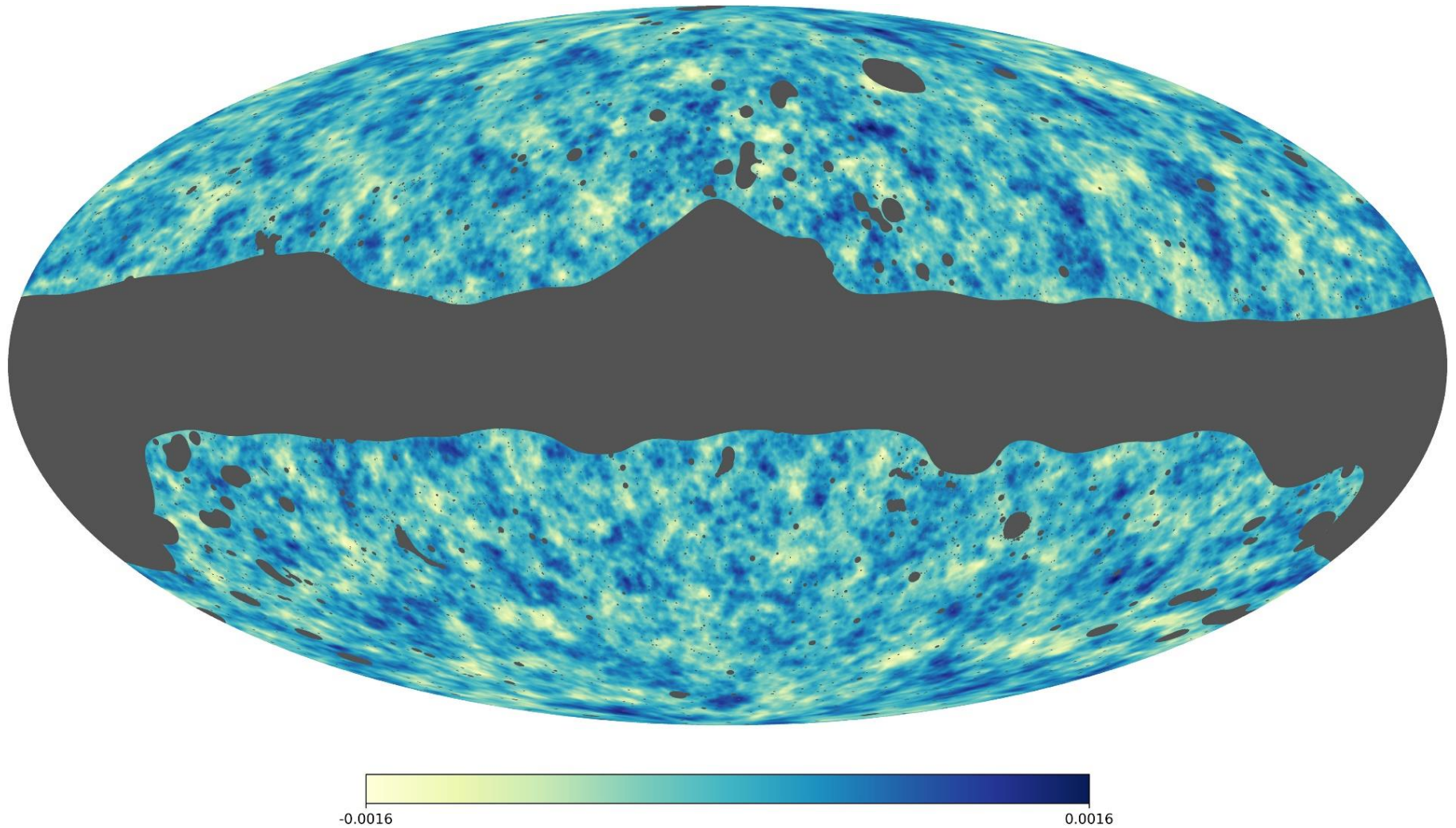
Lensing reconstruction (concept)

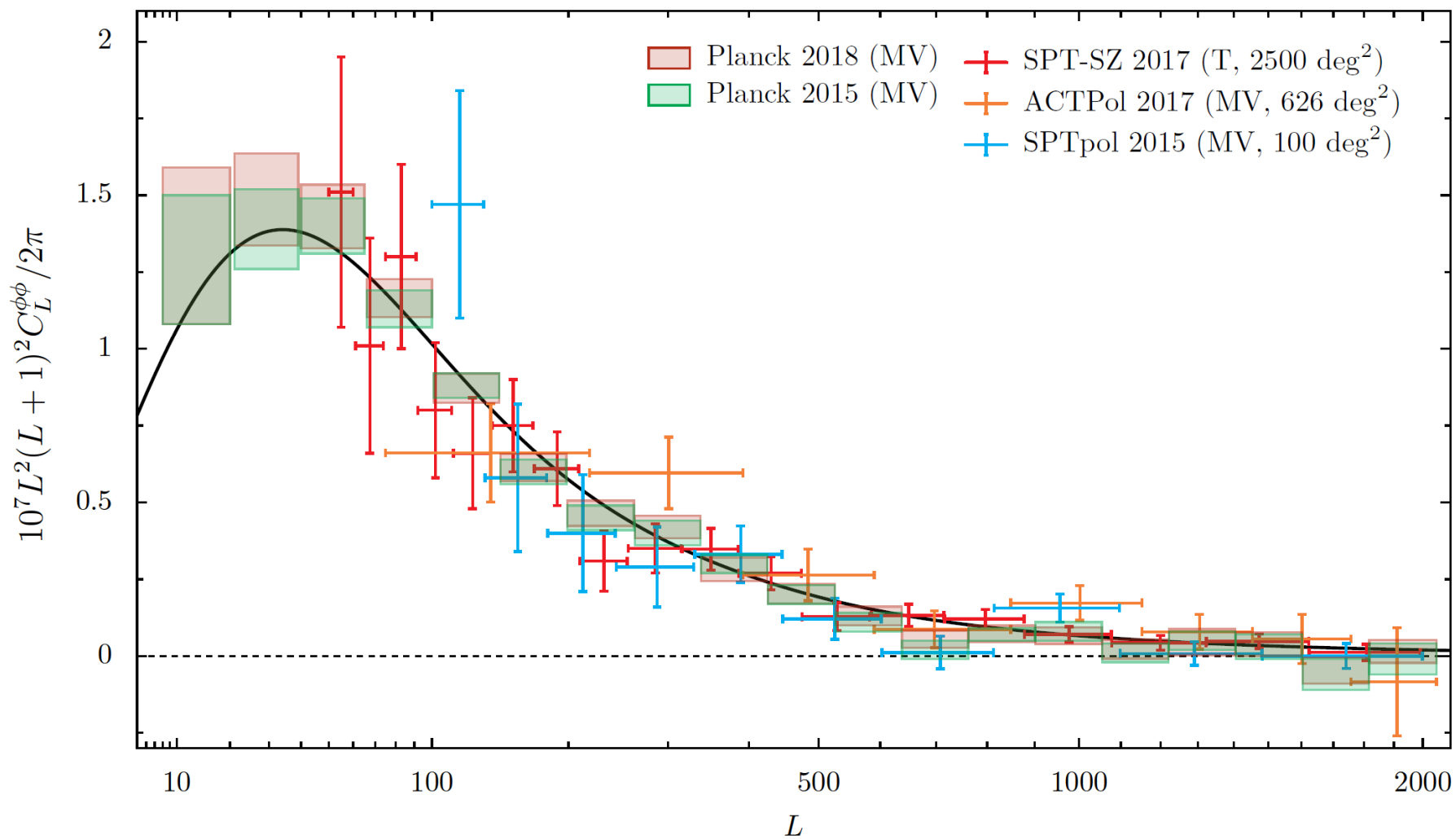


Measure spatial variations in magnification and shear

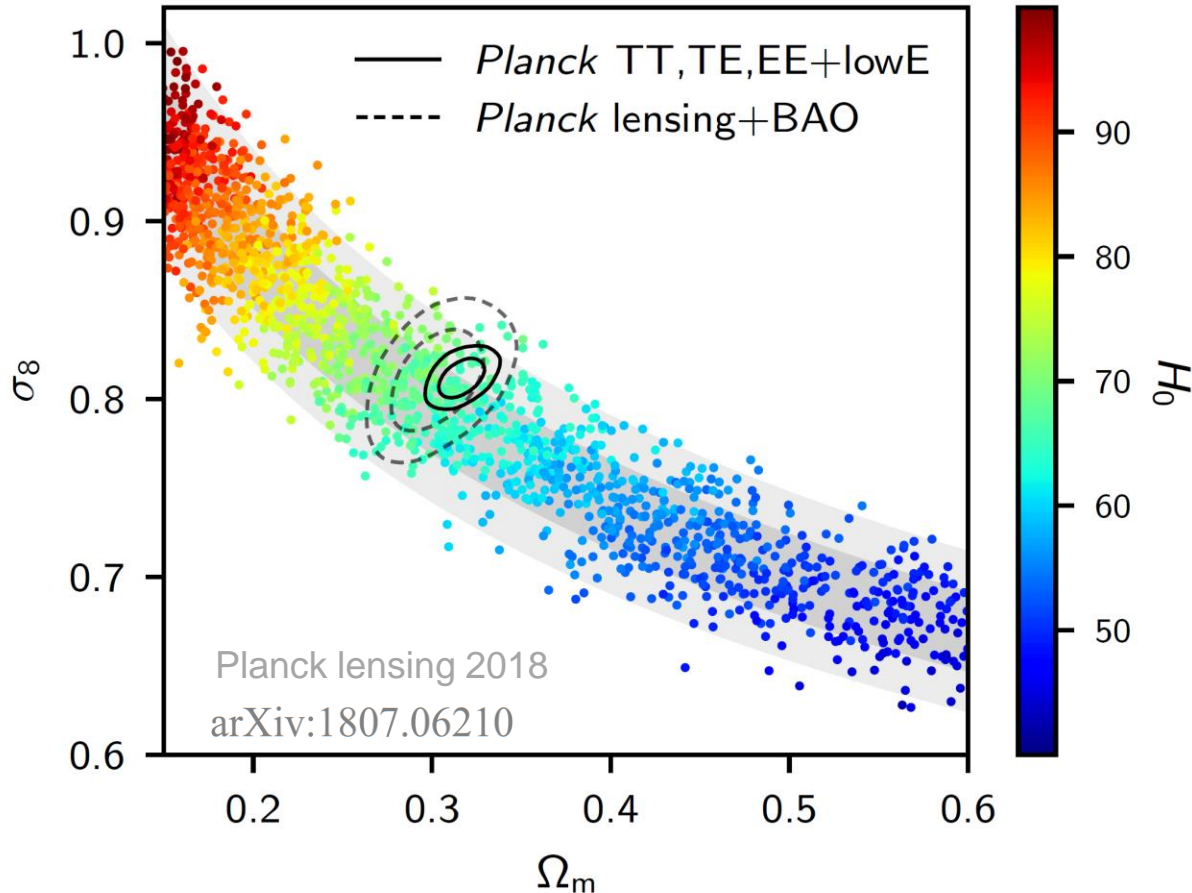
Use assumed unlensed spectrum, and unlensed statistical isotropy

Map of the gradient-mode lensing





Planck CMB lensing parameters



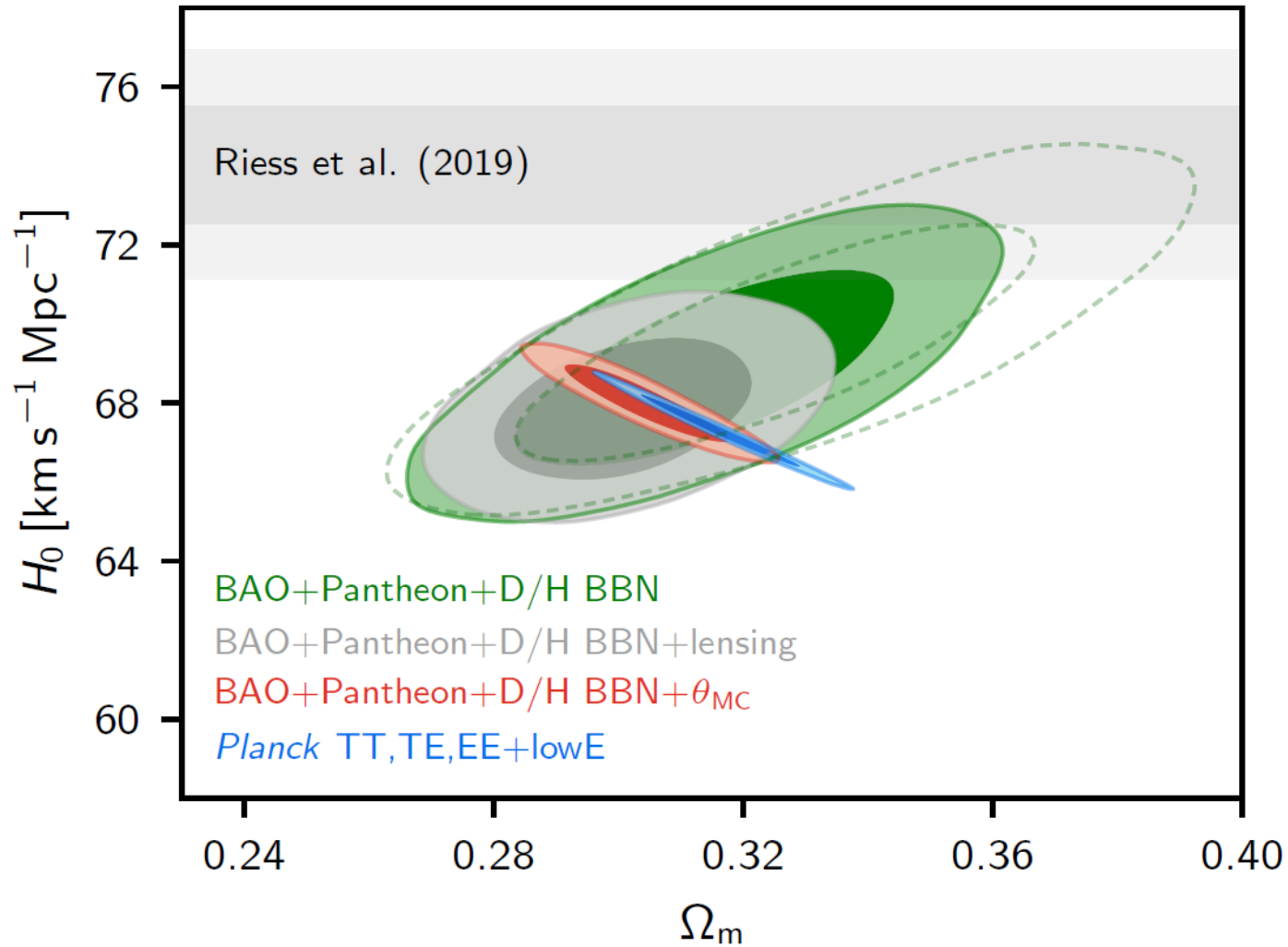
CMB lensing + BAO inverse distance ladder (with $\Omega_b h^2$ prior from abundance measurements)

$$\left. \begin{aligned} H_0 &= 67.9^{+1.2}_{-1.3} \text{ km s}^{-1} \text{ Mpc}^{-1}, \\ \sigma_8 &= 0.811 \pm 0.019, \\ \Omega_m &= 0.303^{+0.016}_{-0.018}, \end{aligned} \right\} 68 \%, \text{ lensing+BAO}$$

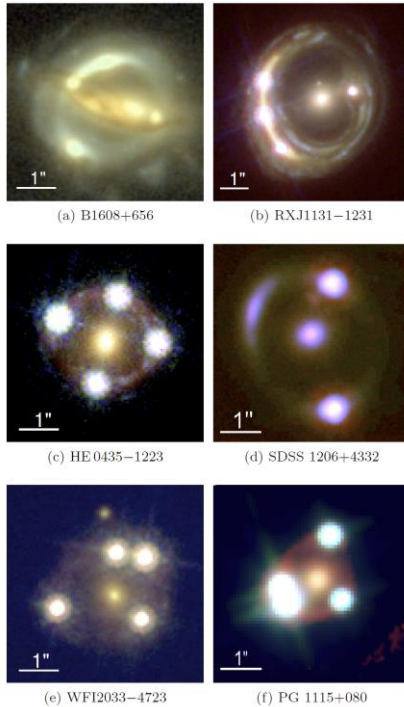
Also adding robust CMB θ_* constraint:
 $H_0 = 68.0 \pm 0.7$ (68 %, lensing+BAO+ θ_*)

("Lensing-only" priors: $\Omega_b h^2 = 0.0222 \pm 0.0005$, $n_s = 0.96 \pm 0.02$, $0.4 < h < 1$)

Independent Λ CDM inverse distance ladder is also consistent with Planck

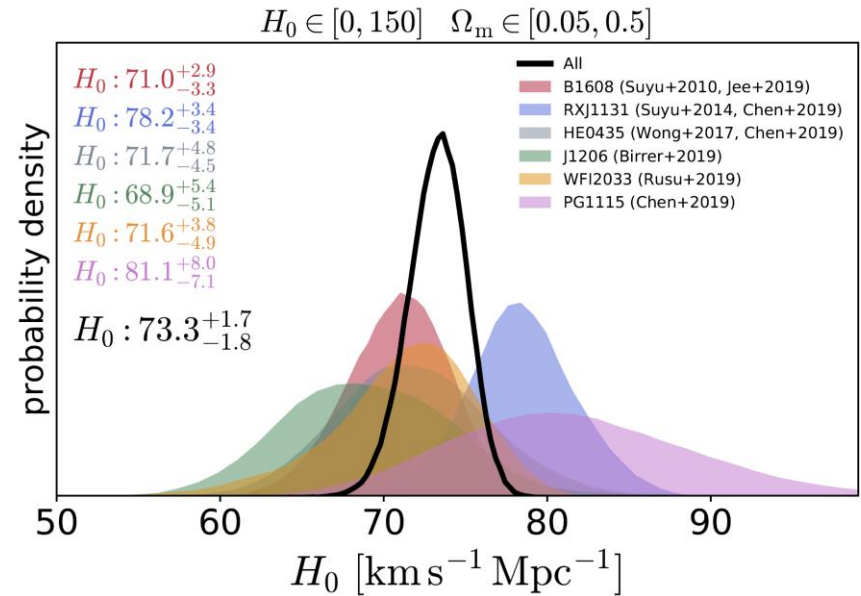


Strong Lensing



Lens modelling etc..

$$D_{\Delta t} \equiv (1 + z_d) \frac{D_d D_s}{D_{ds}}$$



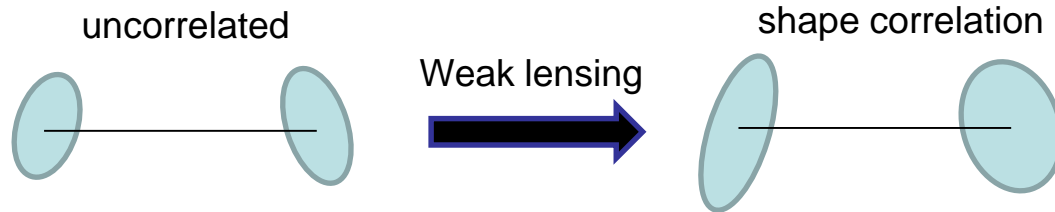
H0LiCOW: $H_0 = 73.3^{+1.7}_{-1.8} \text{ km s}^{-1} \text{ Mpc}^{-1}$

Wong et al. arXiv:1907.04869
(some cosmology dependence)

Independent of CMB and local distance ladder and mostly redshift $z > \sim 0.1$

⇒ tension with CMB independent of very local environment

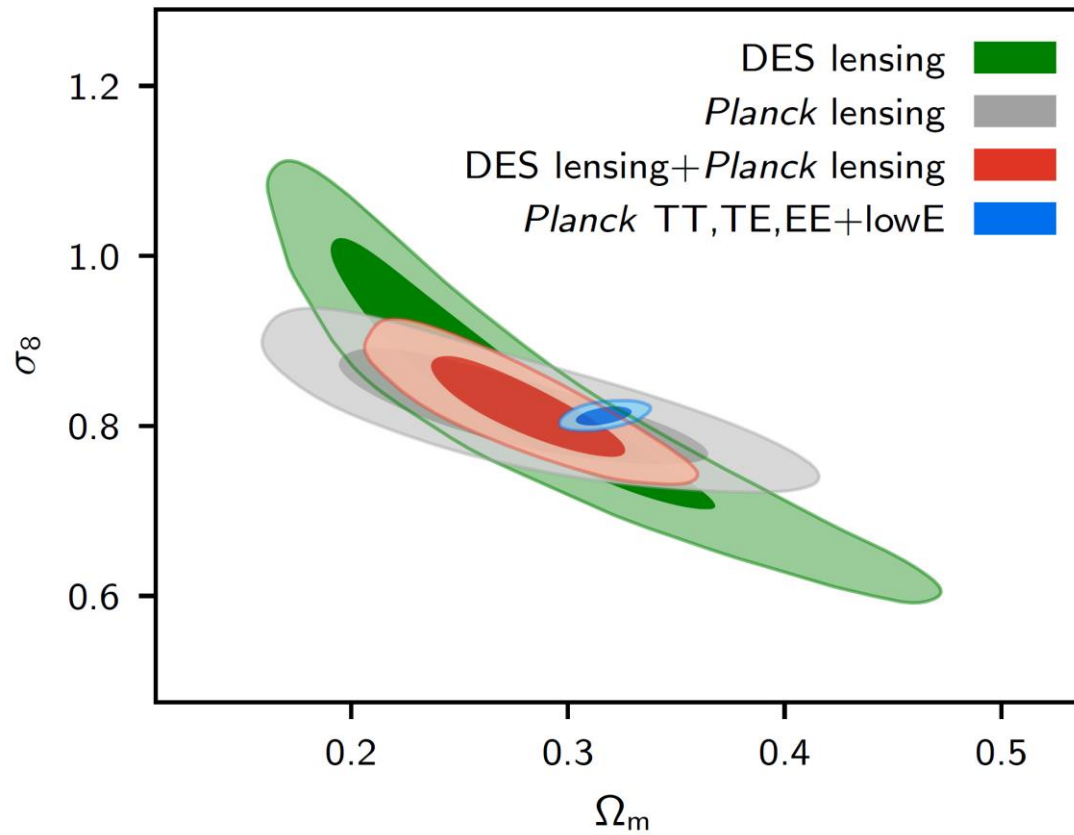
Galaxy weak gravitational lensing – cosmic shear



Potentially clean probe of total matter perturbations and geometry

But, non-linearities, redshift uncertainties, intrinsic alignment, shape biases...

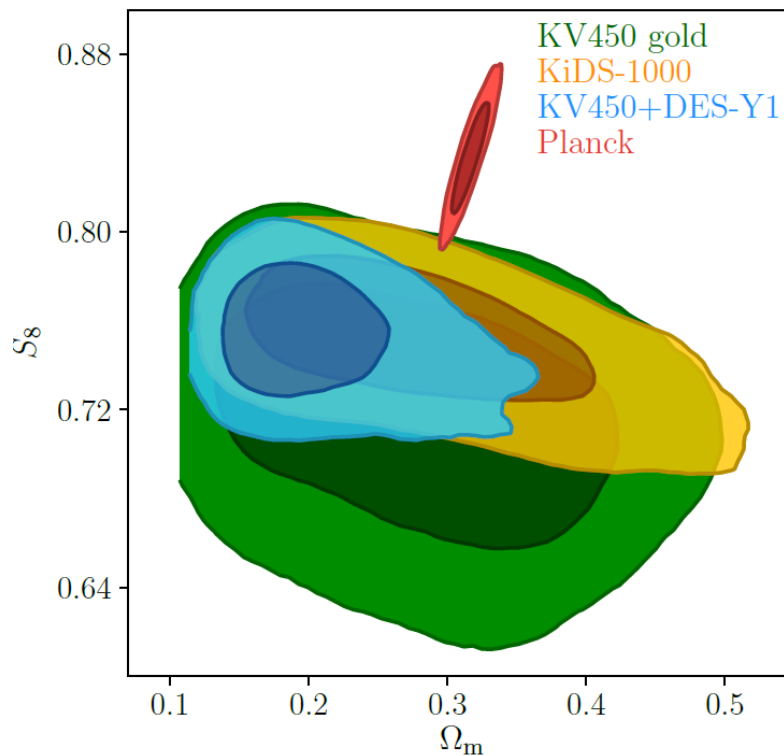
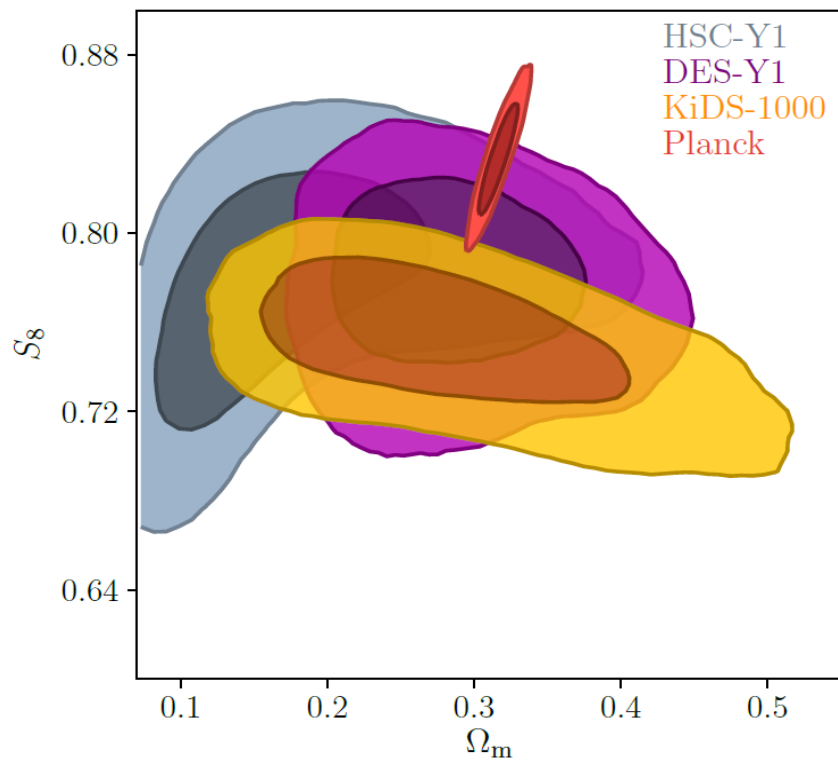
Galaxy and CMB lensing complementary, tighter joint constraints



DES and Planck lensing very consistent

Other galaxy lensing results

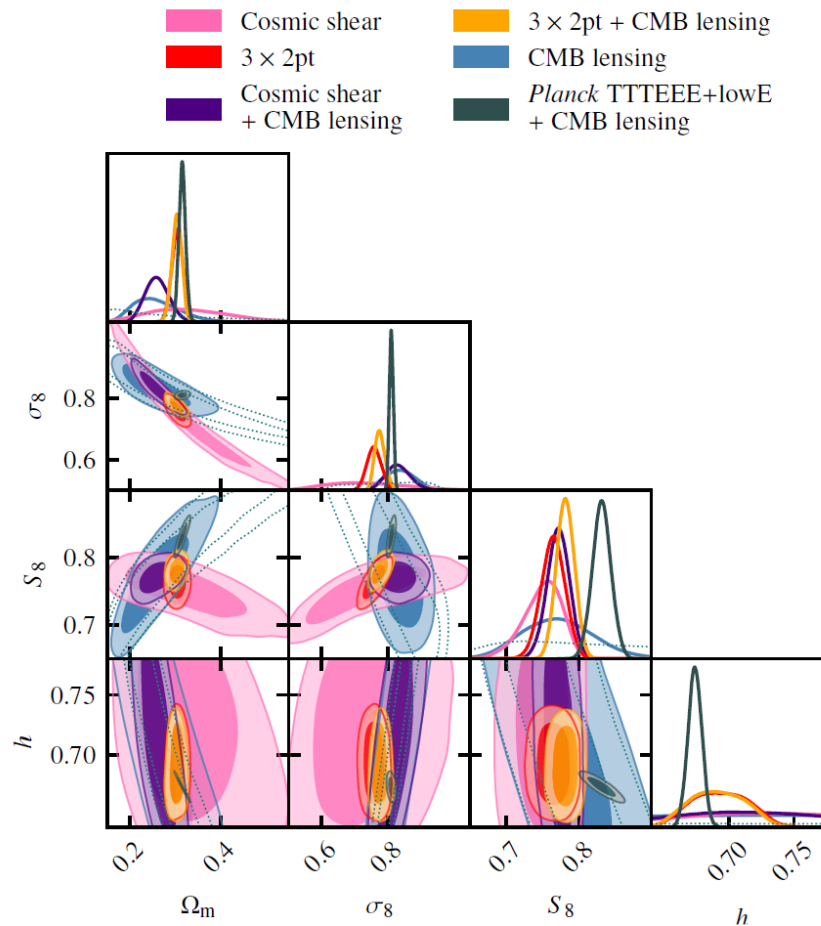
Asgari et al. [arXiv: 2007.15633](https://arxiv.org/abs/2007.15633)



DES+KiDs significant tension with Planck
on $S_8 = \sigma_8 (\Omega_m / 0.3)^{0.5}$

Galaxy lensing + galaxy counts

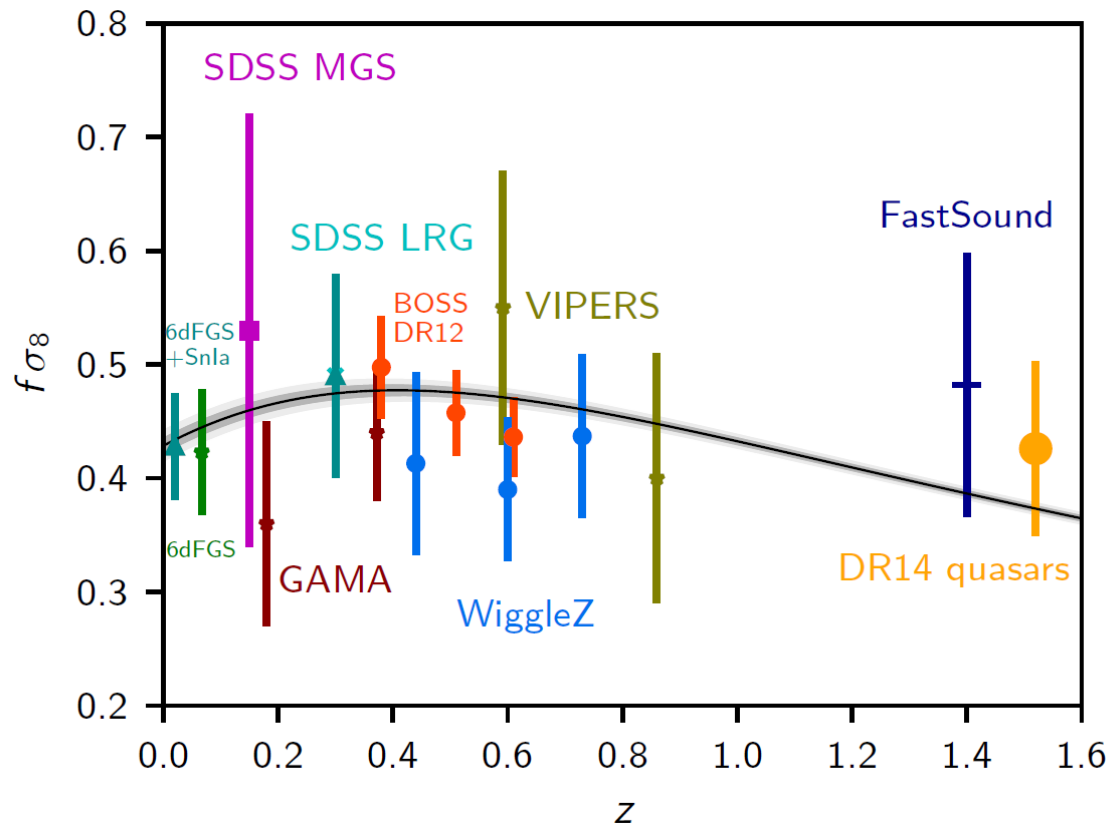
also depends on galaxy bias parameters



KiDS-1000
+BOSS: 3x2pt
arXiv:2010.16416

Marginally consistent/moderate tension with Planck

Redshift Distortions



Currently no compelling evidence for deviations from Planck Λ CDM in LSS observations.

Possible solutions to the H_0 tensions

Biases in data or underestimated error bars

- inverse distance ladder and CMB consistent \Rightarrow *both* CMB and BAO being wrong?
- Local H_0 and strong lensing independent; multiple local distance ladders
but Feedman et al result lower and strong lensing errors relatively large

New physics prior to recombination:

- decrease sound horizon r_d : BAO and Planck H_0 both shift proportionately
- other changes that affect relevant inferred parameters (e.g. $\Omega_m h^2$)

New physics at lower redshift/dark energy/modified gravity

- but $w > -1$ only makes H_0 from Planck *lower*
- have to fit BAO and $H(z)/H_0$ from supernovae (or find problem with supernovae)

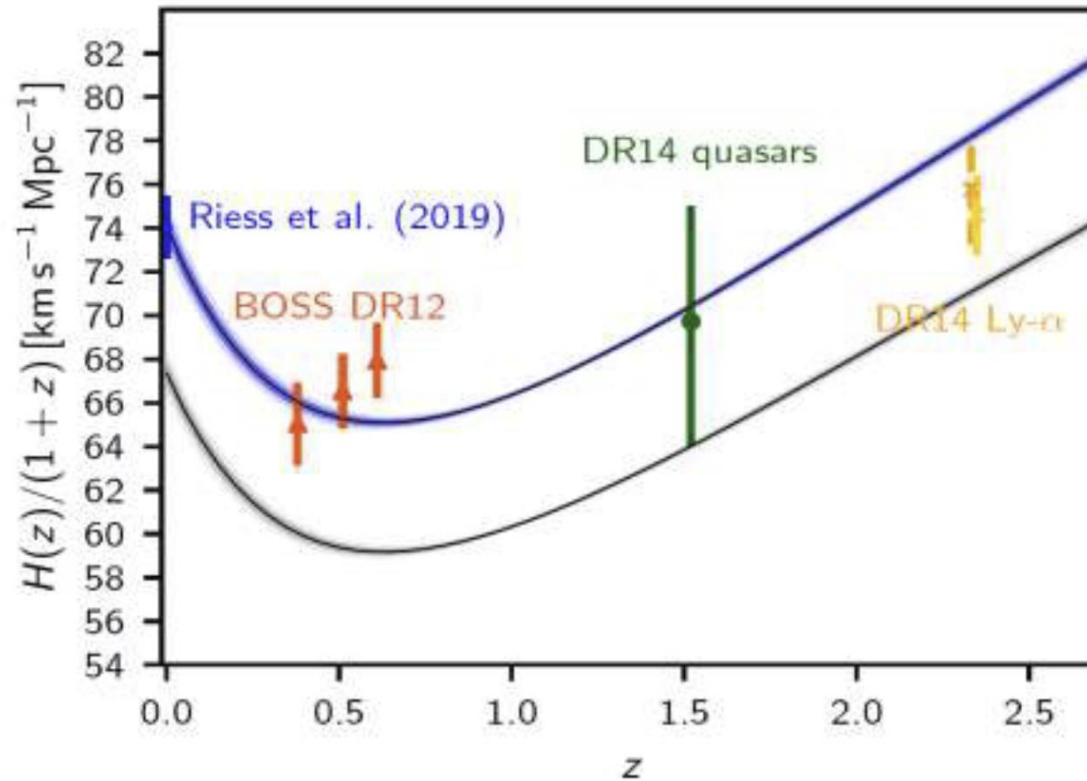
New physics/very unusual conditions in our local neighbourhood

- strong lensing results then in tension?

Some combination of the above

New early universe physics – decrease sound horizon r_d by $O(10\%)$

e.g. increase expansion rate, decrease sound speed, shift recombination, ..



But, simple models e.g. extra relativistic degrees of freedom ($N_{\text{eff}} \neq 3.046$)
not favoured by Planck spectra (and disfavoured by BBN D/H)

*Different models change the CMB spectra in distinctly different ways
e.g. via changes to matter-radiation equality, damping scale, peak phases*

Difficult but not impossible to fit current data

e.g. trade changes from new physics with changes in $\Omega_c h^2, \Omega_b h^2, A_s, n_s, \dots$)

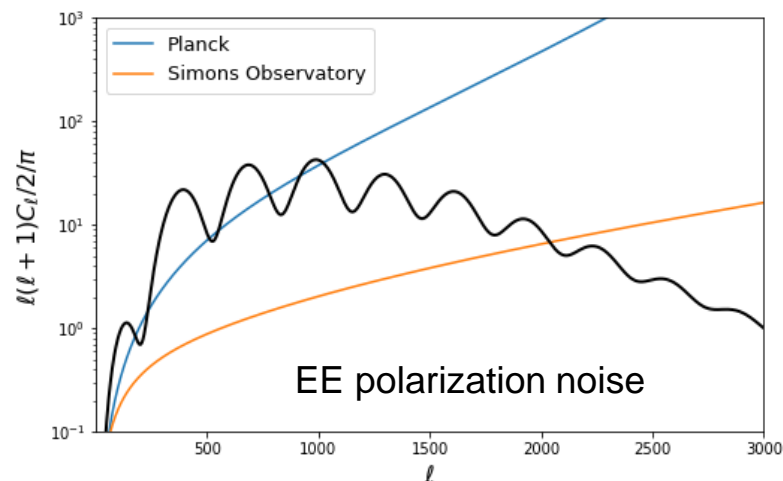
⇒ Almost impossible to *also* fit Λ CDM polarization to cosmic variance

⇒ new “easily” detectable EE/TE signal that does not fit Λ CDM

(note: if new physics is the solution, current Λ CDM measurements of parameters likely to be significantly wrong, e.g. significant implications for inflation n_s)



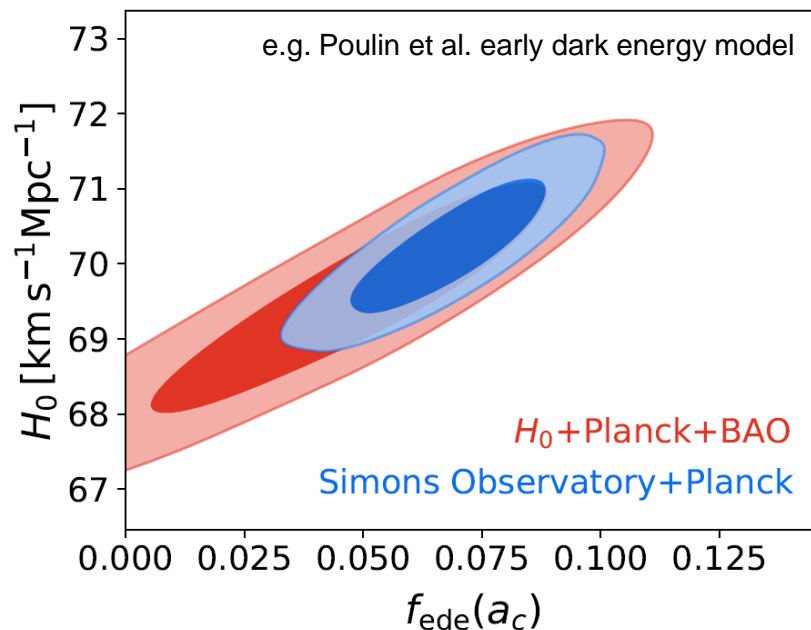
High resolution/sensitivity polarization:
precision small-scale EE, TE, TT power spectrum



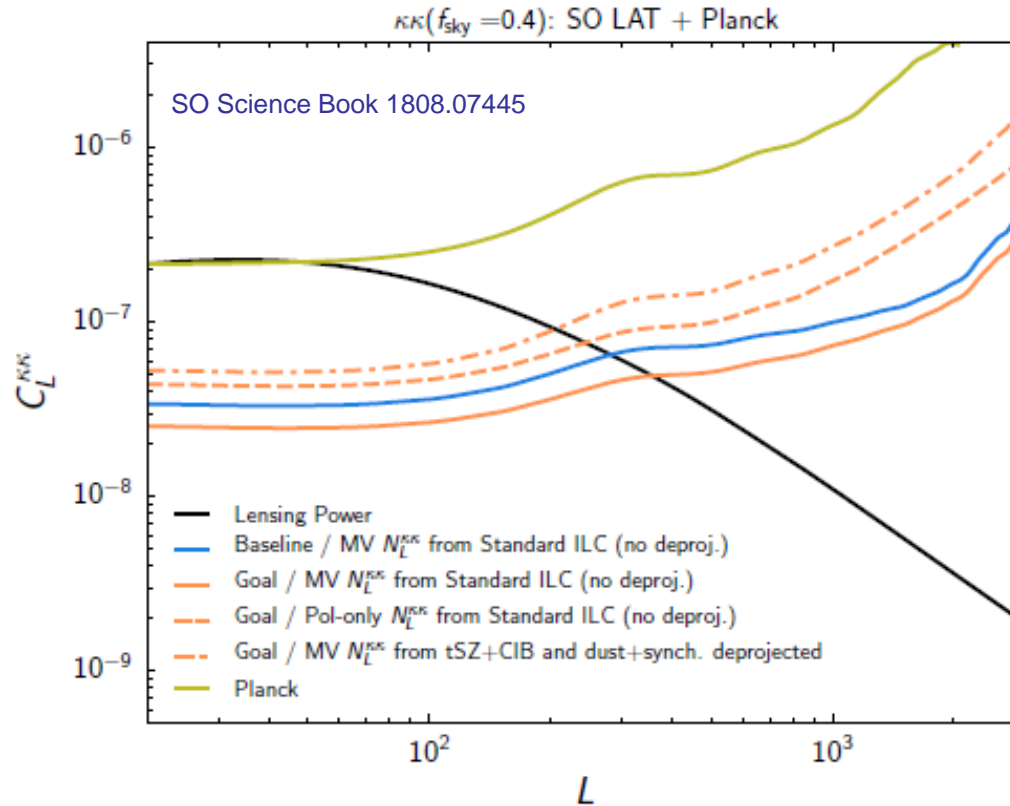
+ ActPol, SPTpol (now)
+ CMB-S4 (beyond)

If $H_0 > 71 \text{ km s}^{-1} \text{Mpc}^{-1}$,
new pre-recombination physics
likely detectable at $> 5\sigma$ soon

*Distinct physical models give
different precision predictions*



Simons Observatory CMB Lensing forecast



CMB lensing constraints will dramatically improve
... and completely different systematics to galaxy lensing

+ new galaxy lensing from LSST, EUCLID, etc.

Conclusions

Λ CDM concordance between CMB, BAO, SN, CMB lensing, BBN (except lithium)

... and *BAO and CMB are the cleanest and most robust probes*

H_0 tension 1-5+ σ

- Complex indirect measurements, but multiple independent or semi-independent probes
- New pre-recombination physics at 5-10 % level “easily” detectable soon with CMB polarization

- can test *reason for discrepancy* \Rightarrow distinguish new physics

No models currently attempted are compelling or great fits.

- *and why does it look in so many ways just like Λ CDM?*

CMB lensing and galaxy lensing complementary and very different systematics

Some tensions in late-time σ_8 measurements, but complex and evolving

- More powerful LSS measurements soon could give clearer indication
(*unless statistical power all soaked up by nuisance parameters in the complex modelling*)

Cobaya: Code for Bayesian Analysis

Jesus Torrado, AL

[arXiv:2005.05290](https://arxiv.org/abs/2005.05290)

Python parameter sampling framework: likelihoods -> parameter MC samples

Optimizations to exploit different speeds of multiple dependent theory/likelihood modules each with multiple nuisance parameters

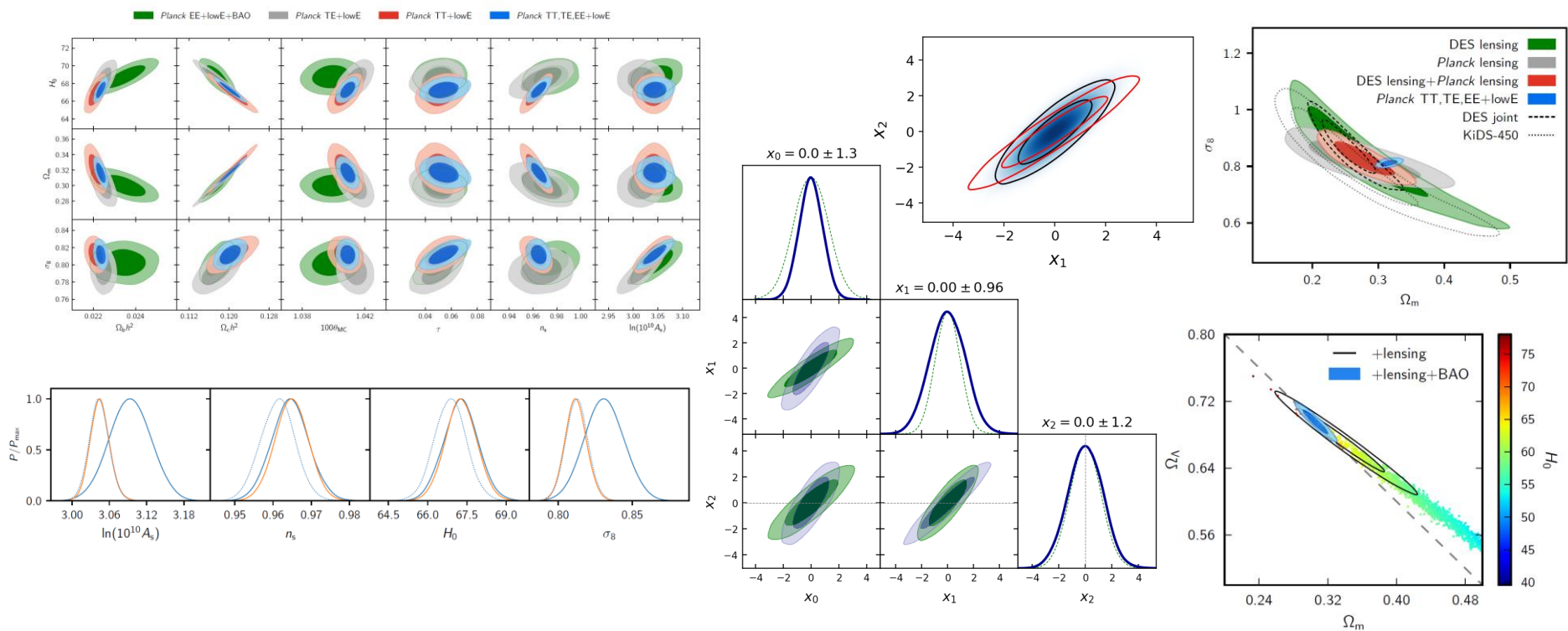
<https://github.com/CobayaSampler/cobaya>

<https://cobaya.readthedocs.io/>



GetDist 1.0: Python Monte Carlo Sample Analyser

<https://getdist.readthedocs.io> (arXiv:1910.13970)



+ interactive GUI, KDE, PCA, convergence, latex, tables